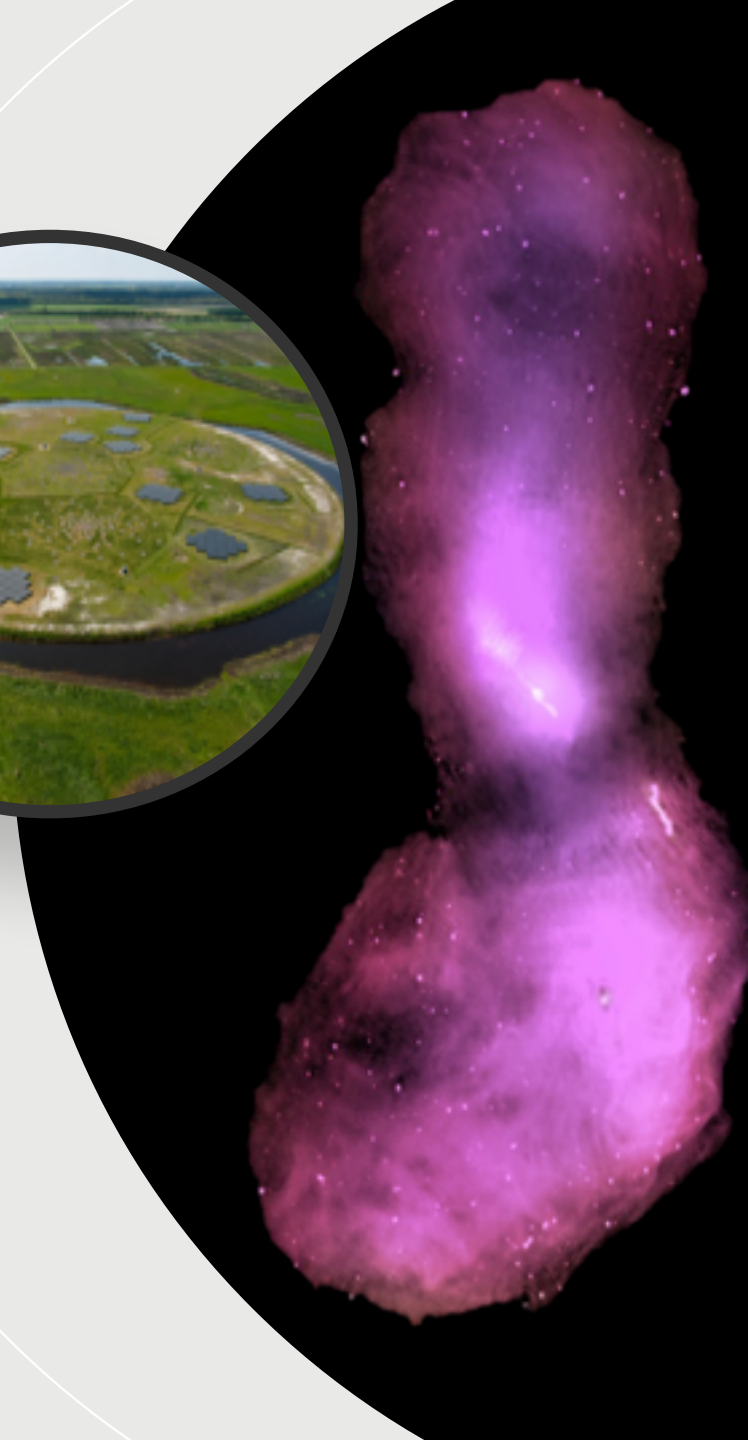
ASTRON

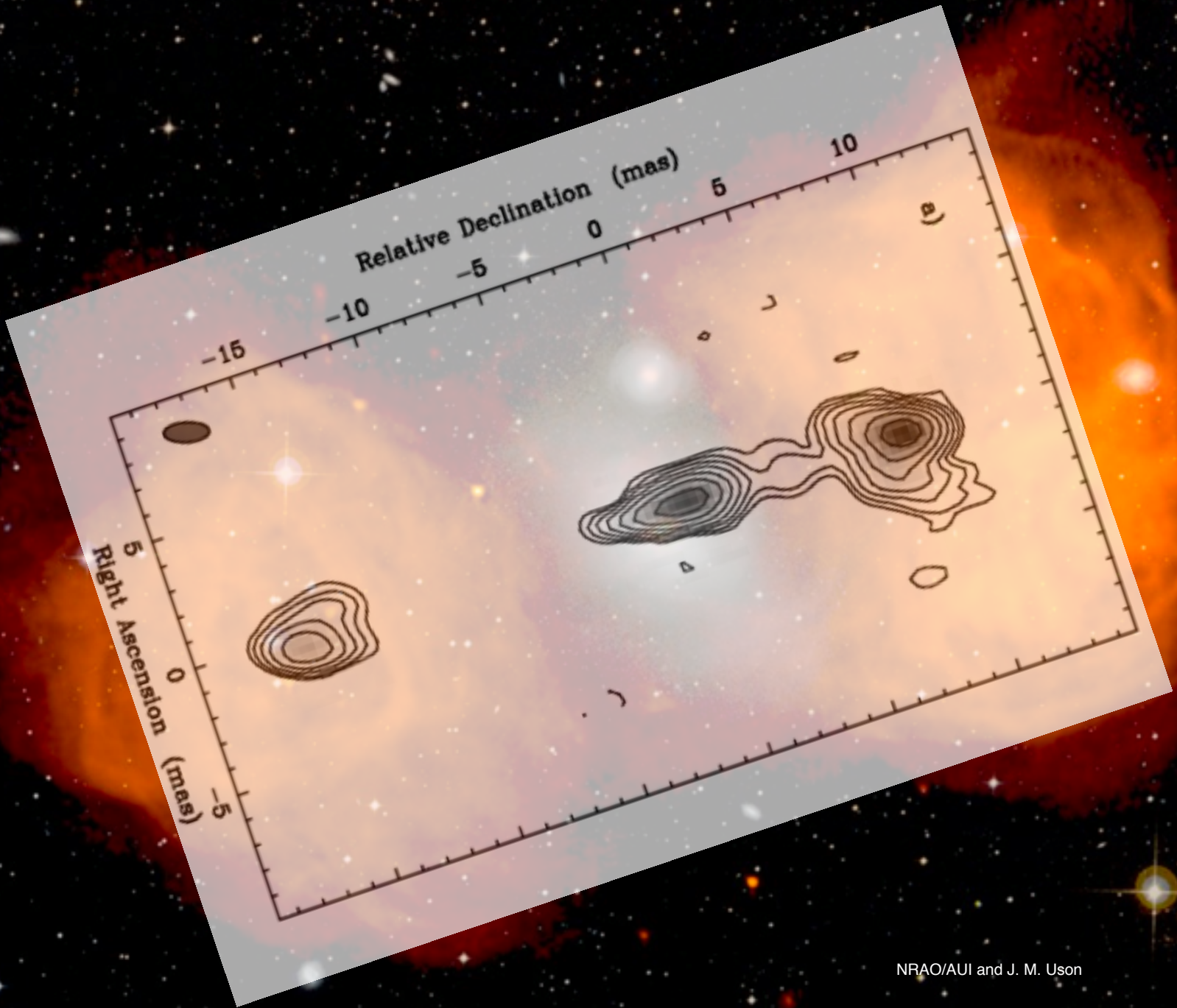
Why are there so many the peaked spectrum sources in the Universe?



Joe Callingham
ASTRON Postdoctoral Fellow

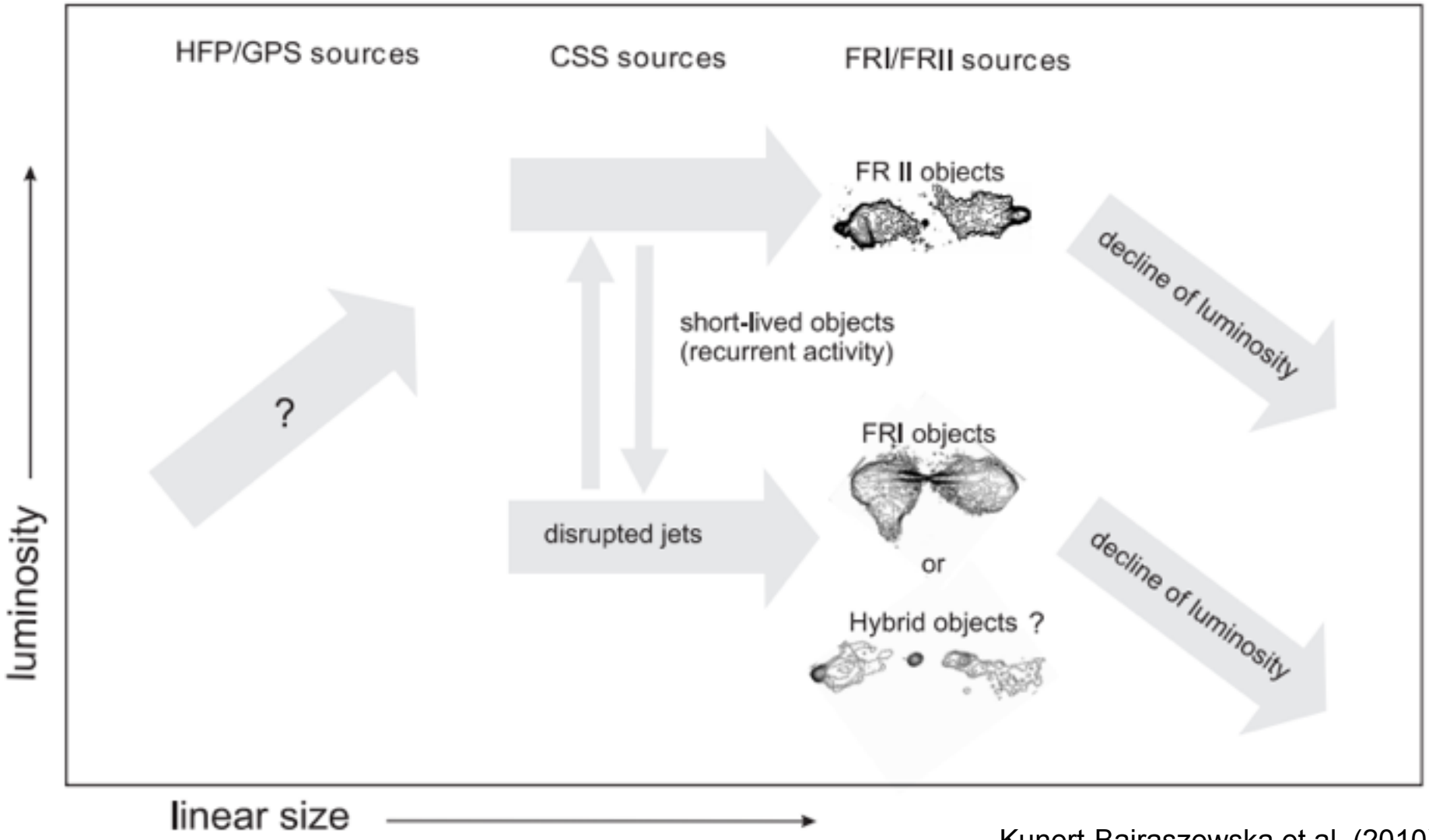
SALFIV,
Sydney, Australia
13th of December 2017

LOFAR



NRAO/AUI and J. M. Uson

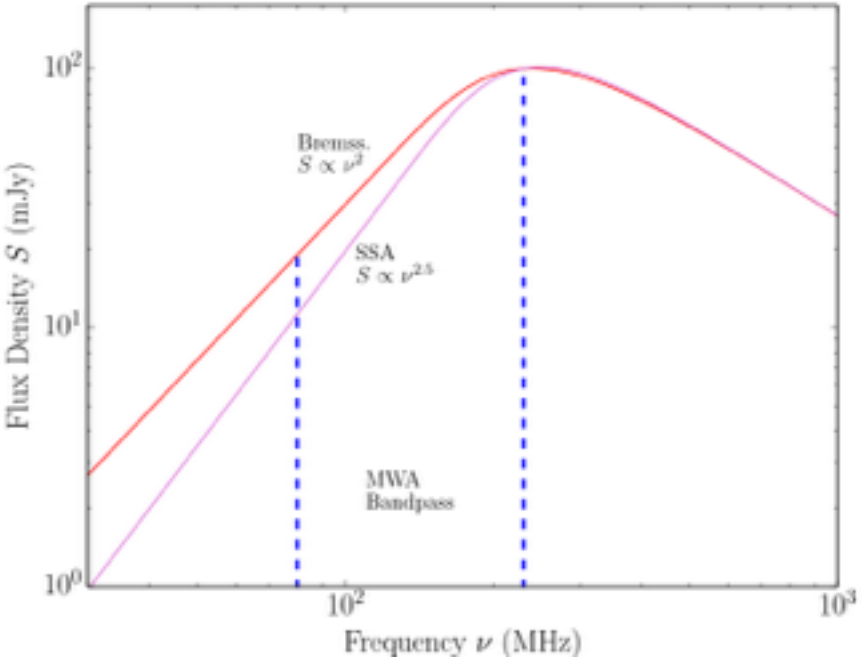
Possible Evolutionary Picture

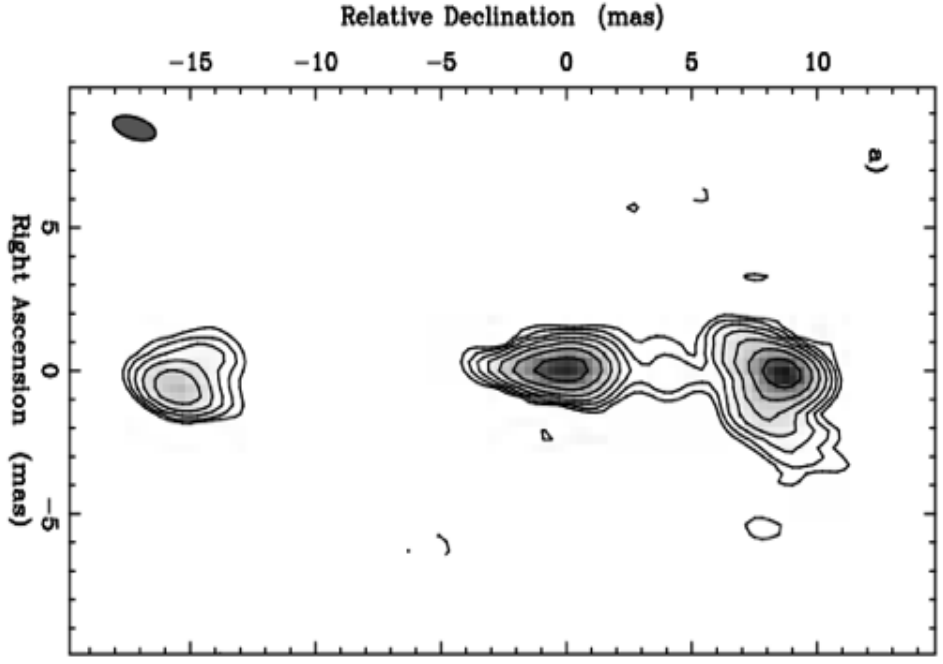


Kunert-Bajraszewska et al. (2010)

What are GPS/CSS Sources?

- GPS = gigahertz-peaked spectrum ; CSS = compact steep spectrum
  - powerful AGN with ***concave*** radio spectra
  - GPS turnover ~ 1 GHz ; CSS turnover ~ 150 MHz (?)
  - small physical sizes: GPS < 1 kpc, CSS ~ 1 – 10 kpc
  - hosts vary: quasars, radio galaxies, and Seyferts





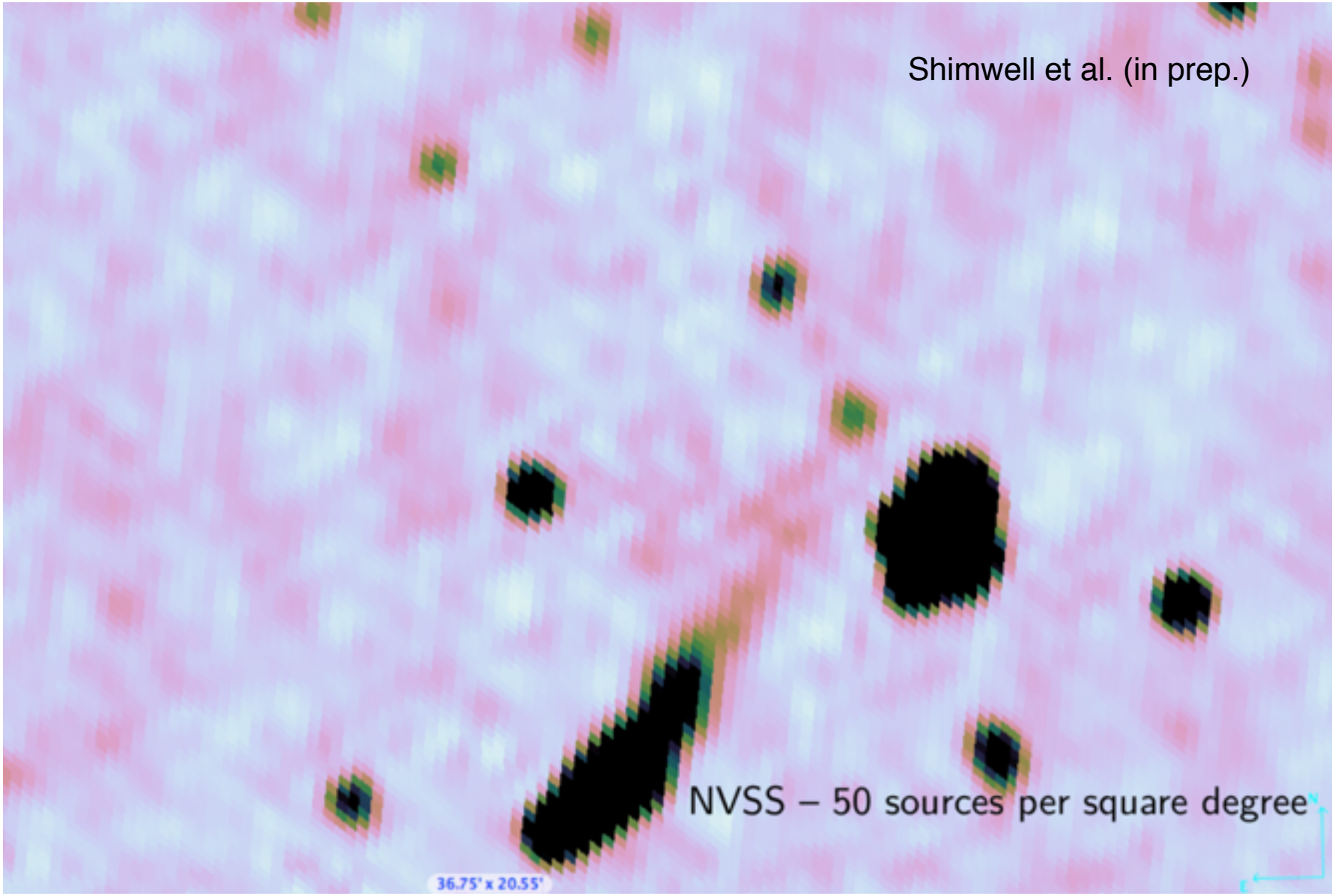
0710+439 (Owsianik et al. 1998)

Tier-1 Survey and Hetdex

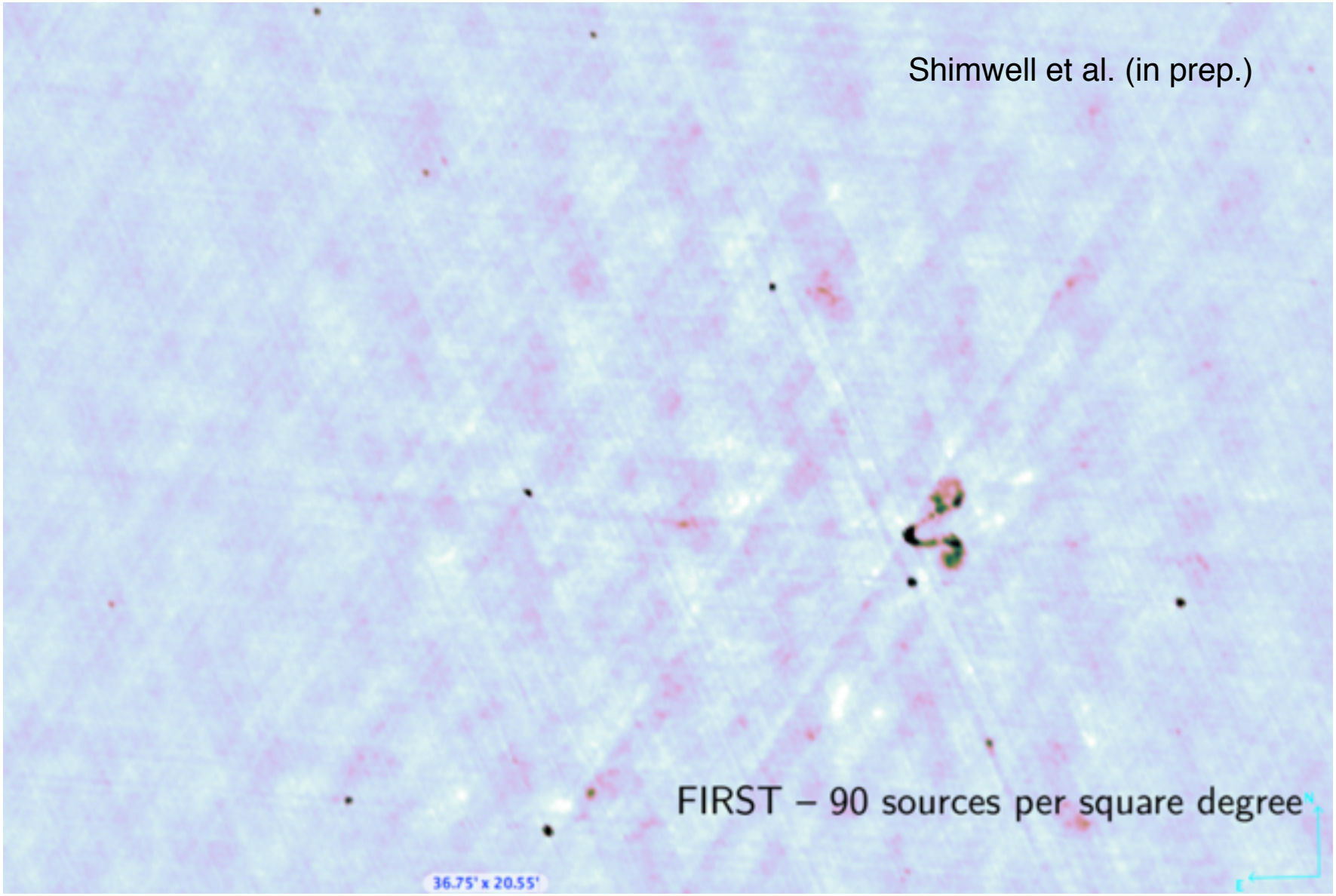
- ~6” resolution
- ~100 μJy/beam rms noise
- ~325,000 sources
- Hetdex release in ~April 2017 under Shimwell et al. (2017)

Shimwell et al. (in prep.)

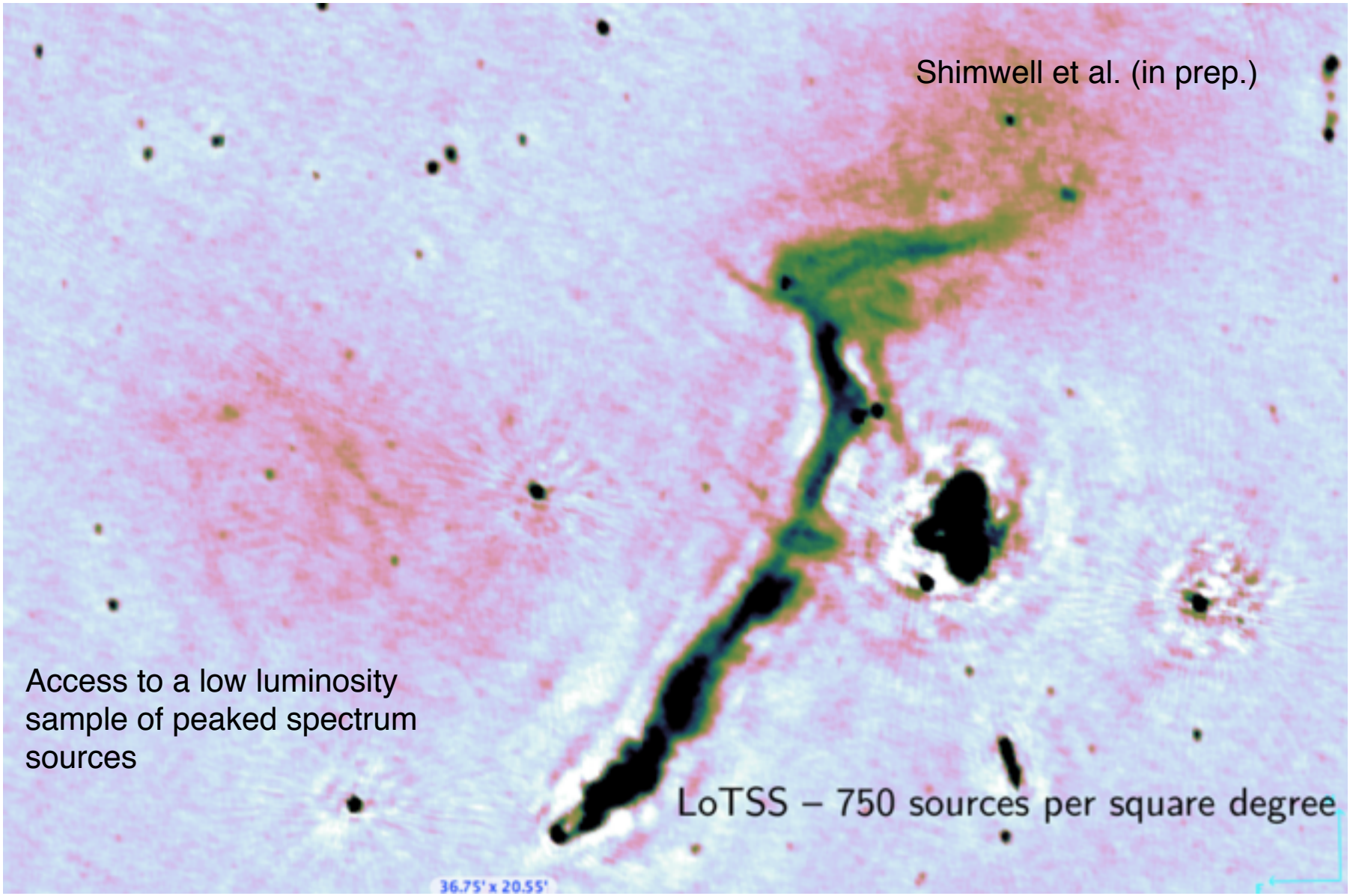
LoTSS


Shimwell et al. (in prep.)
NVSS – 50 sources per square degree
36.75' x 20.55'

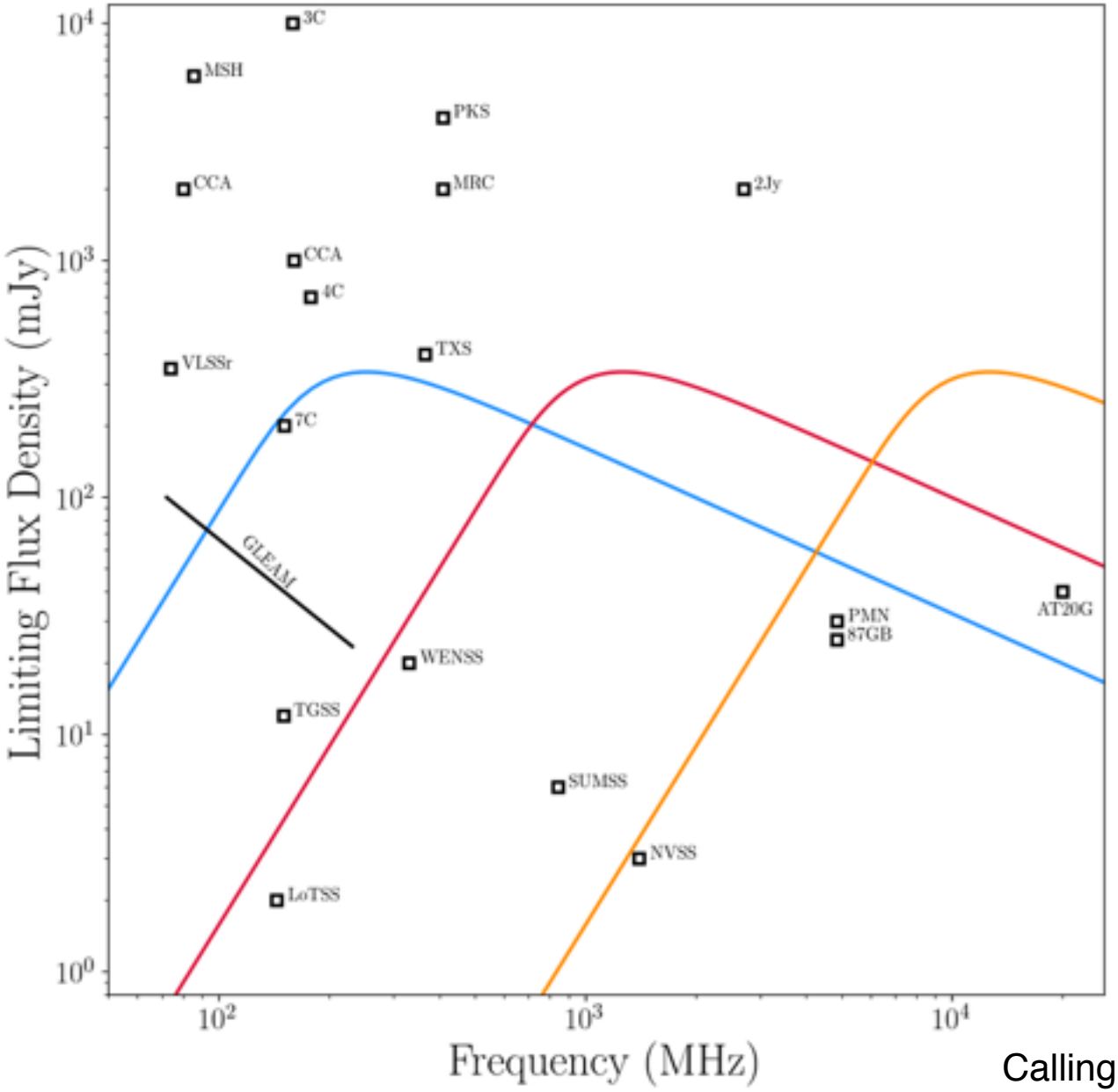
LoTSS


Shimwell et al. (in prep.)
FIRST – 90 sources per square degree
36.75' x 20.55'

LoTSS

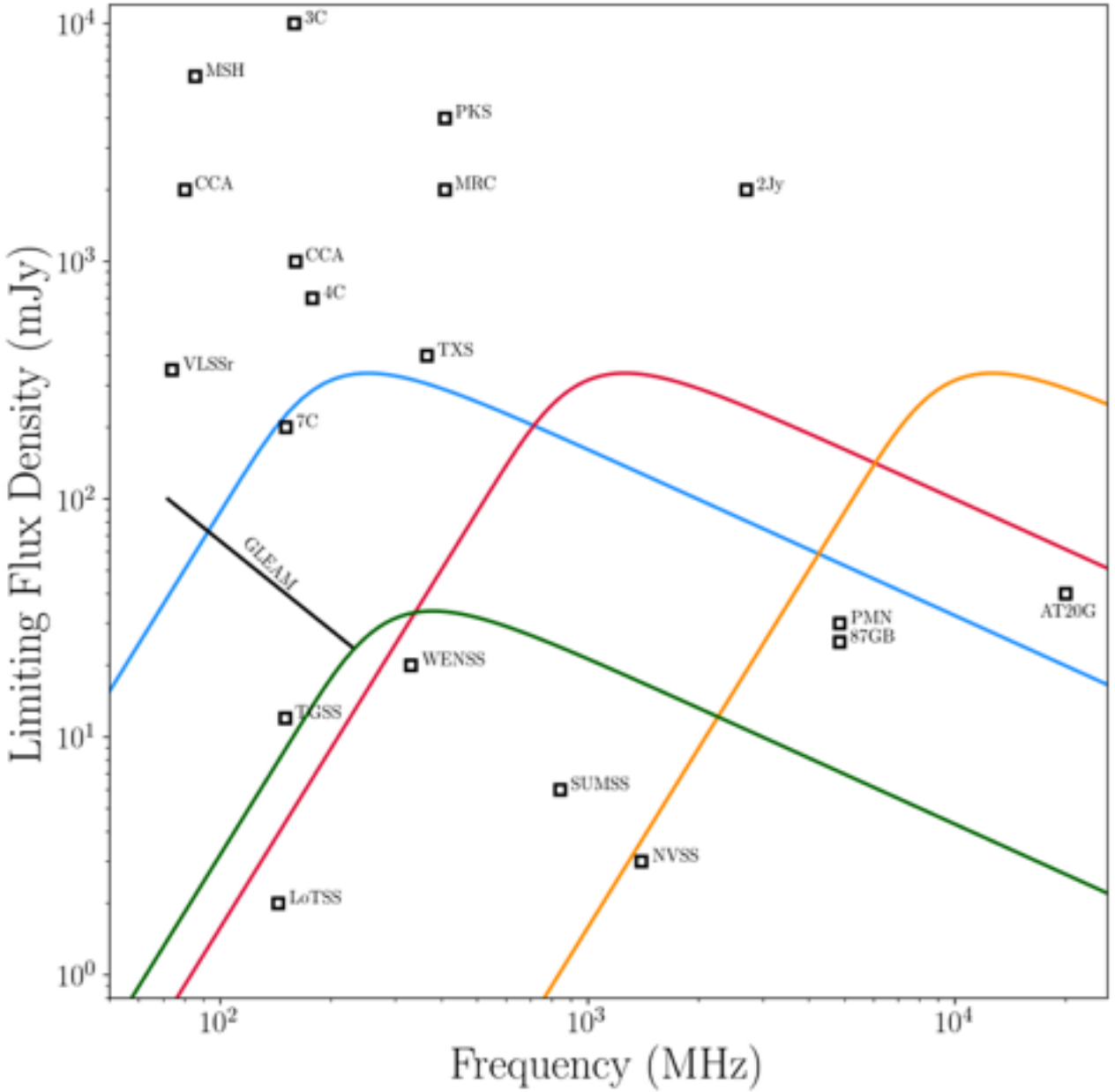

Shimwell et al. (in prep.)
Access to a low luminosity sample of peaked spectrum sources
LoTSS – 750 sources per square degree
36.75' x 20.55'

Which GPS?

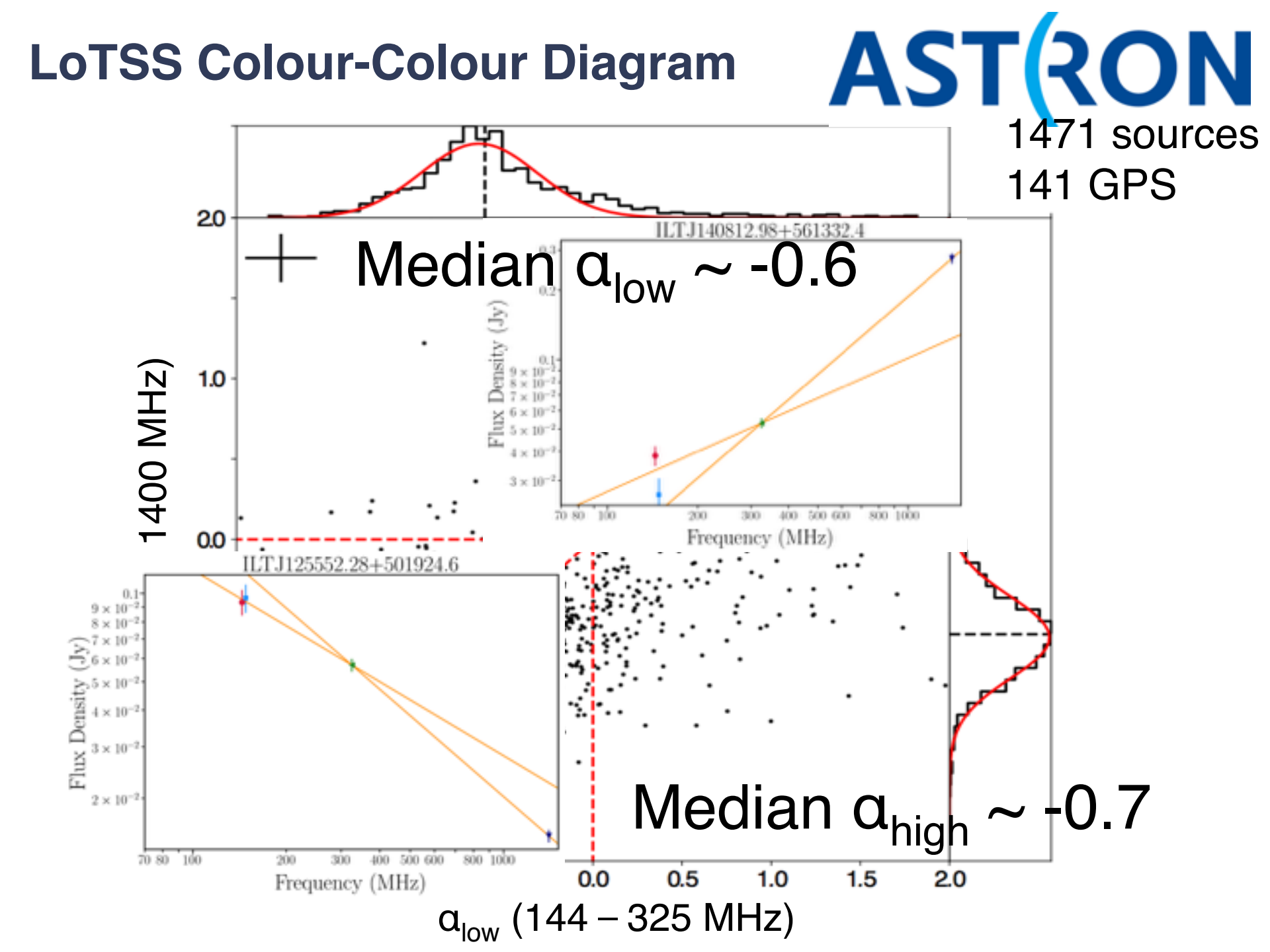


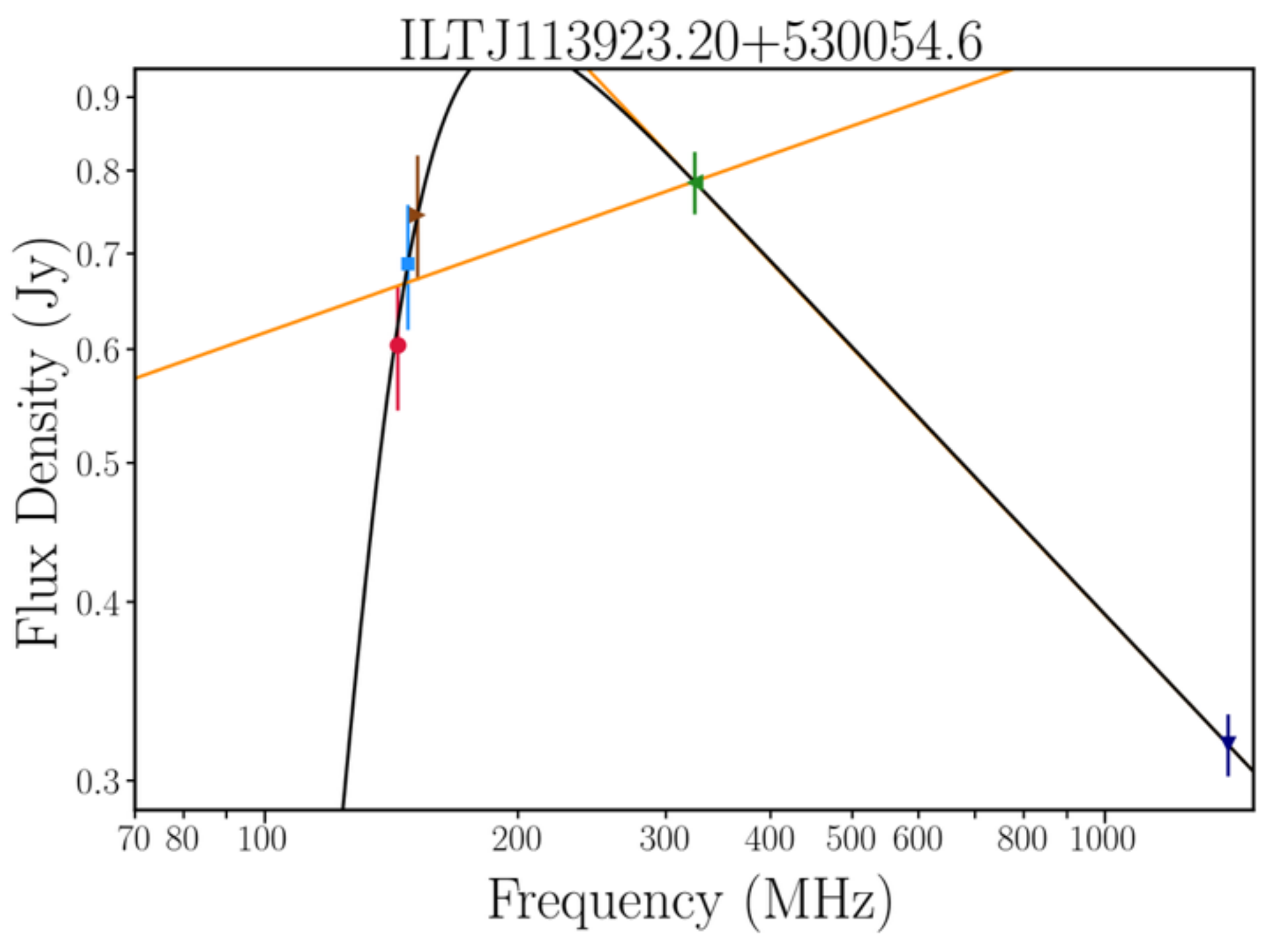
Callingham et al. (2017)

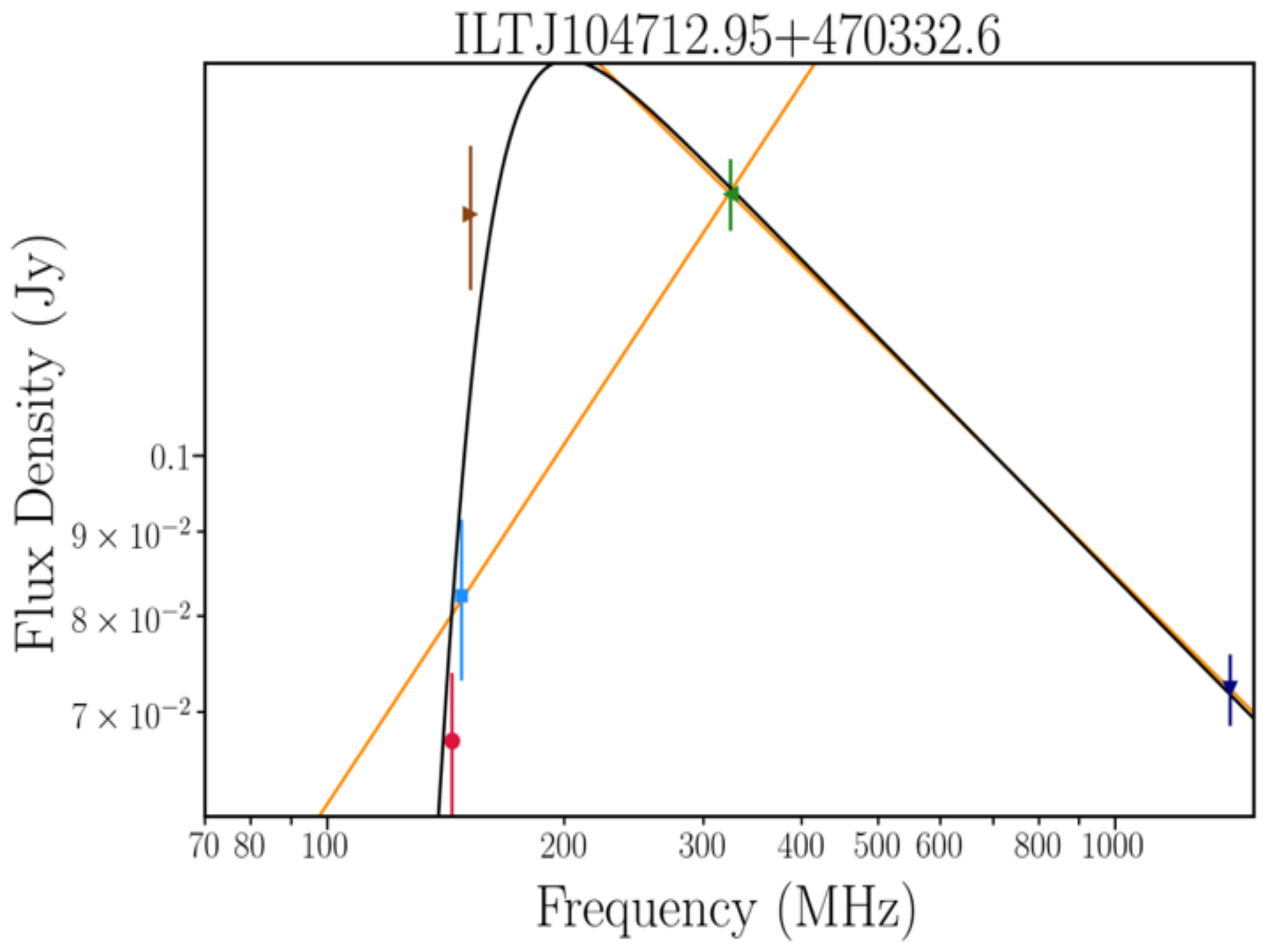
Which GPS?

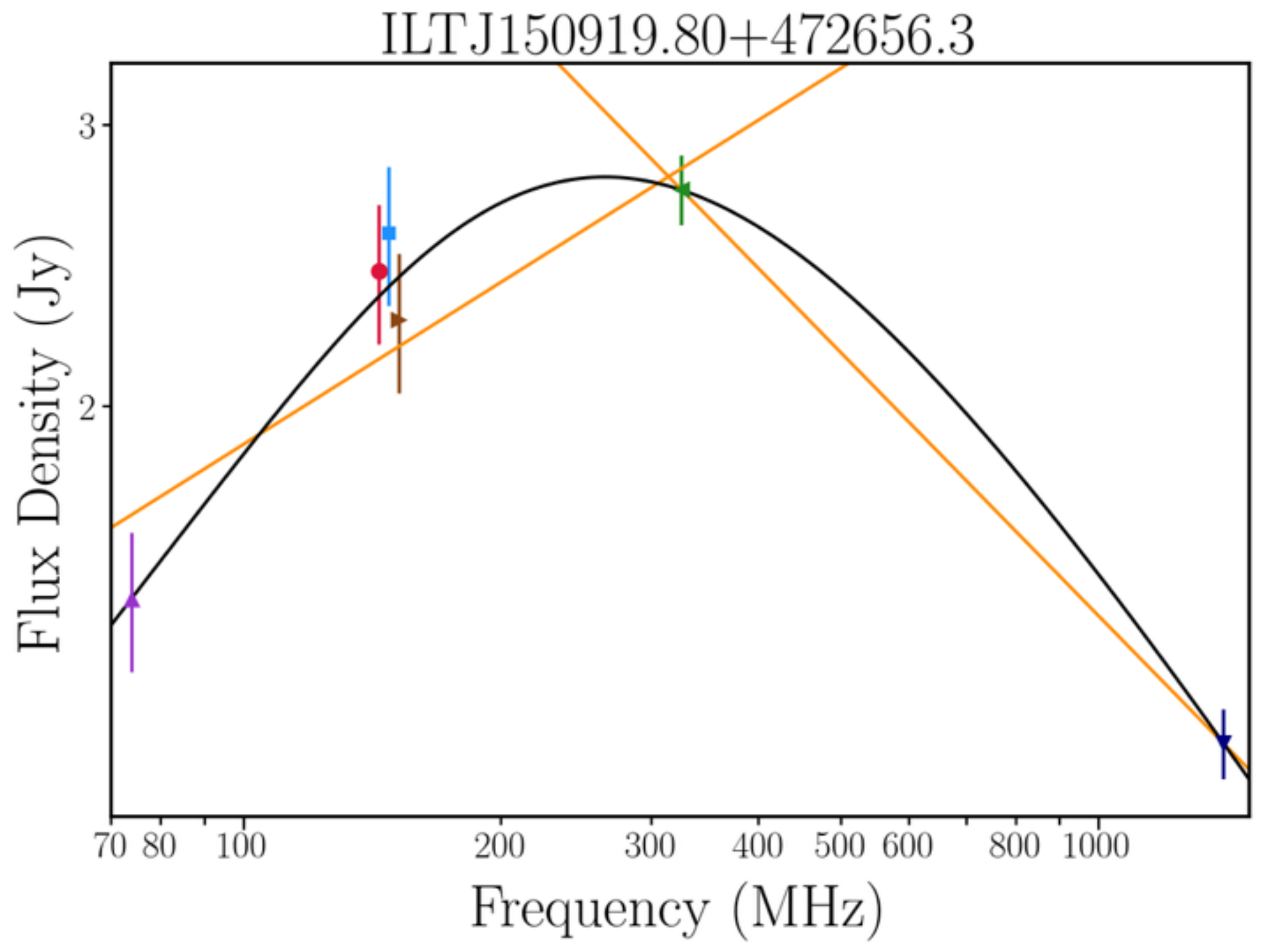

3C
MSH
PKS
CCA
MRC
2Jy
CCA
4C
VLSSr
TXS
7C
GLEAM
WENSS
PMN
87GB
AT20G
TGSS
SUMSS
NVSS
LoTSS
Limiting Flux Density (mJy)
Frequency (MHz)
10^0
10^1
10^2
10^3
10^4
10^2
10^3
10^4

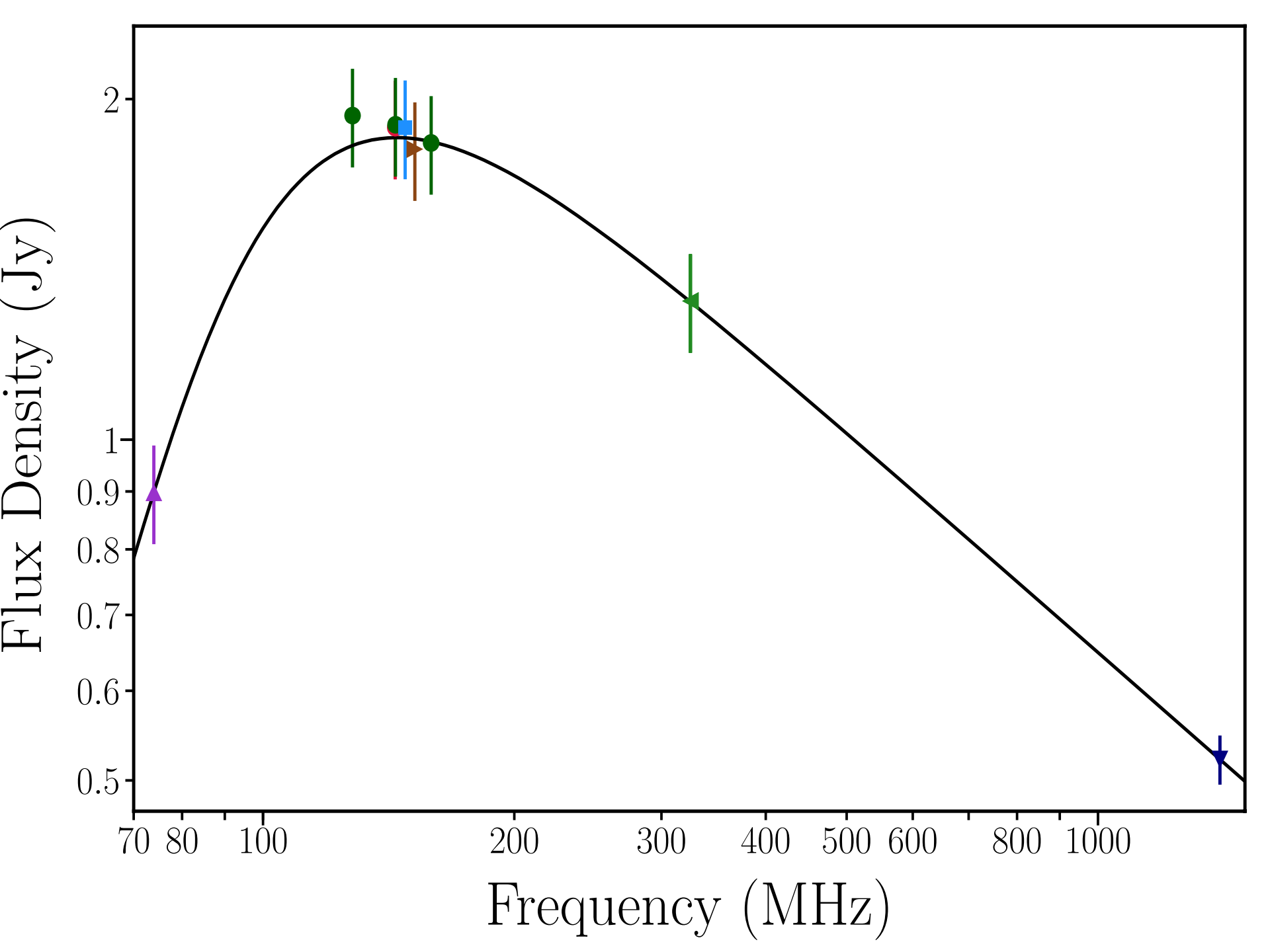
LoTSS Colour-Colour Diagram




ILTJ113923.20+530054.6
Flux Density (Jy)
0.9
0.8
0.7
0.6
0.5
0.4
0.3
70
80
100
200
300
400
500
600
800
1000
Frequency (MHz)

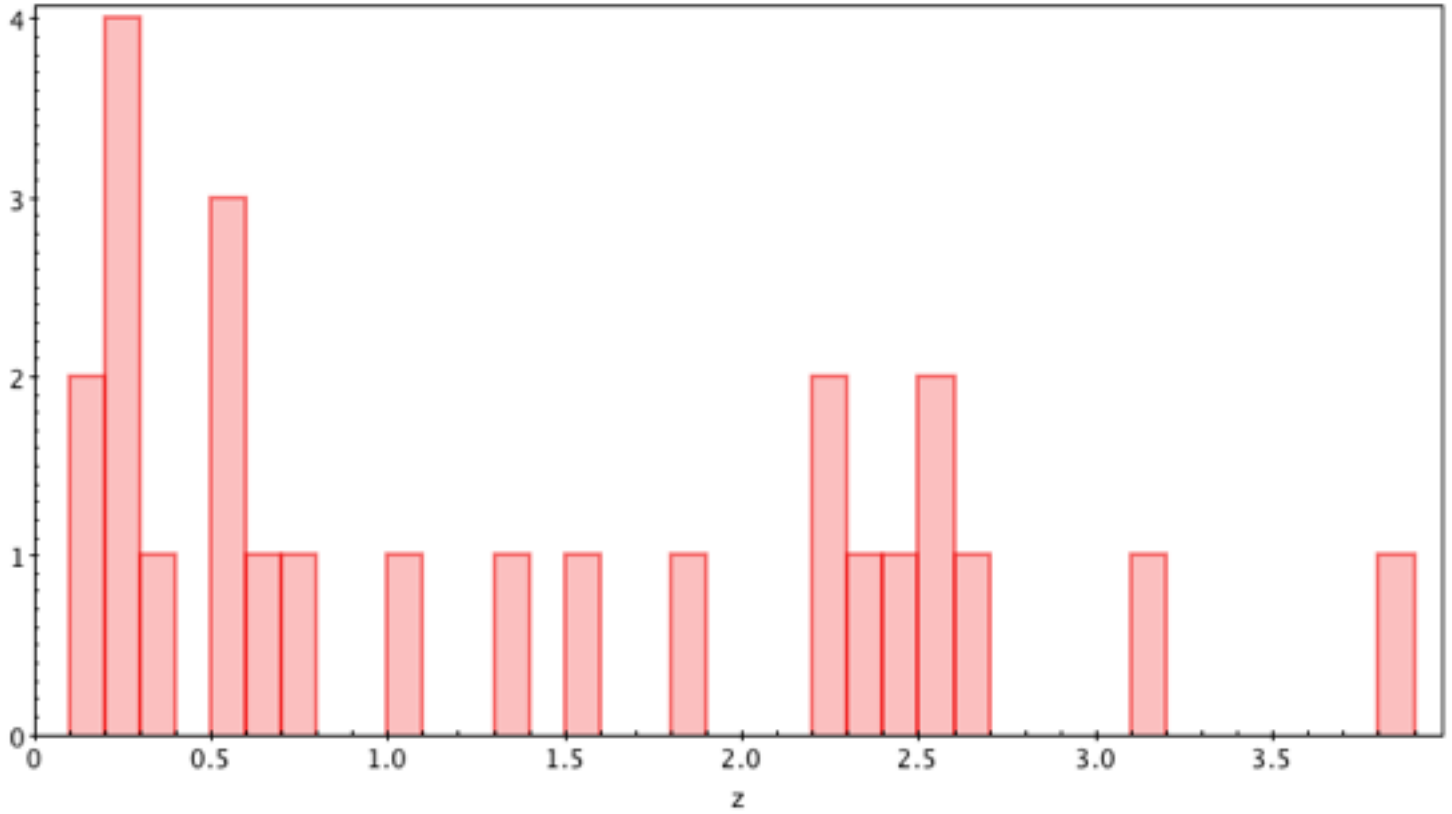

ILTJ104712.95+470332.6
Flux Density (Jy)
0.1
9×10^{-2}
8×10^{-2}
7×10^{-2}
70 80 100 200 300 400 500 600 800 1000
Frequency (MHz)


ILTJ150919.80+472656.3
Flux Density (Jy)
3
2
70
80
100
200
300
400
500
600
800
1000
Frequency (MHz)

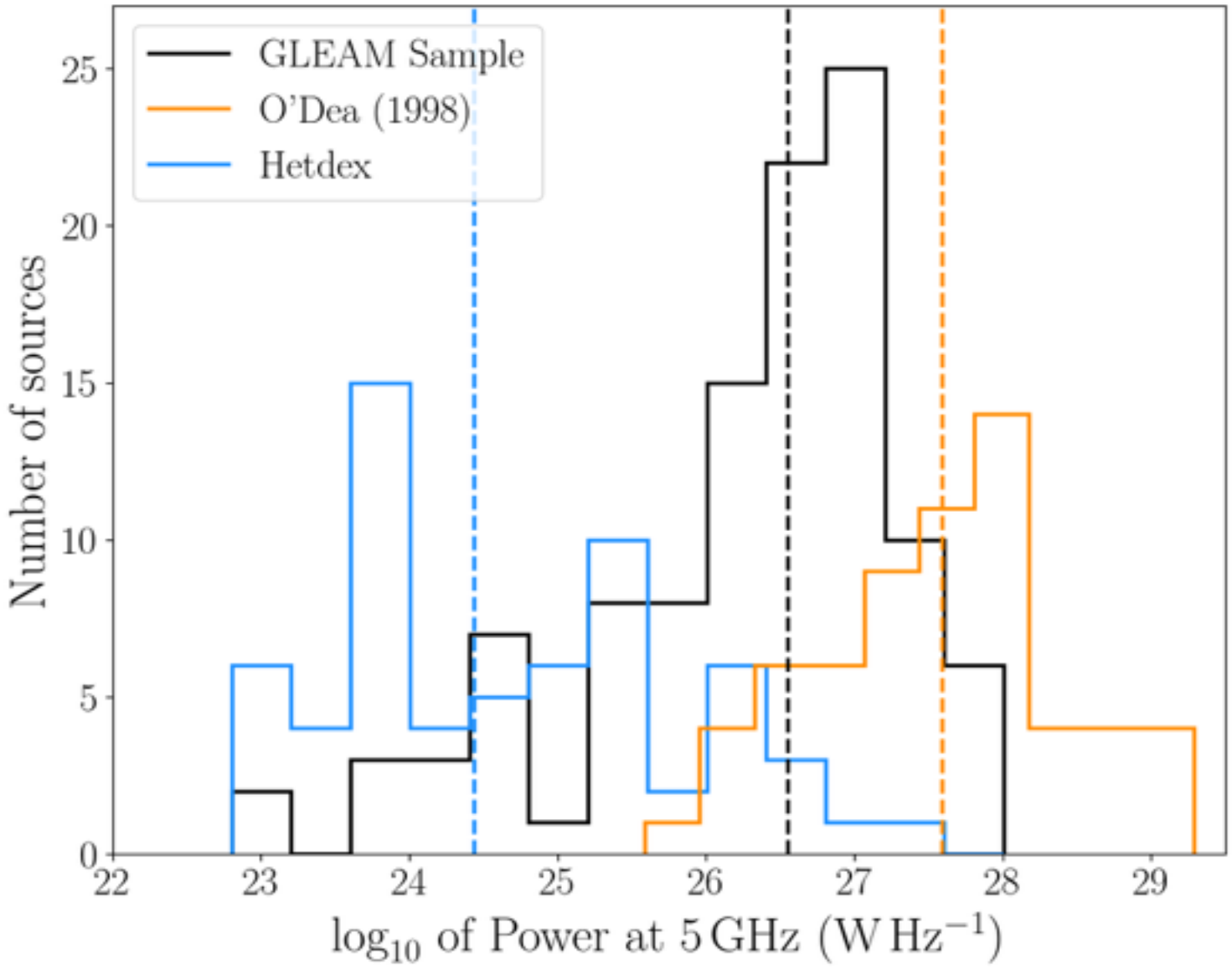

Flux Density (Jy)
2
1
0.9
0.8
0.7
0.6
0.5
70
80
100
200
300
400
500
600
800
1000
Frequency (MHz)

Too many…

- ~10 % of sources that have NVSS/WENSS counterpart are peaked-spectrum
- Obviously completeness issues but compared to complete sample with GLEAM, we have double the number of sources selected at the same frequency (~4.5%). Why?
- 25 of 144 have spectroscopic redshift (SDSS + literature/NED)



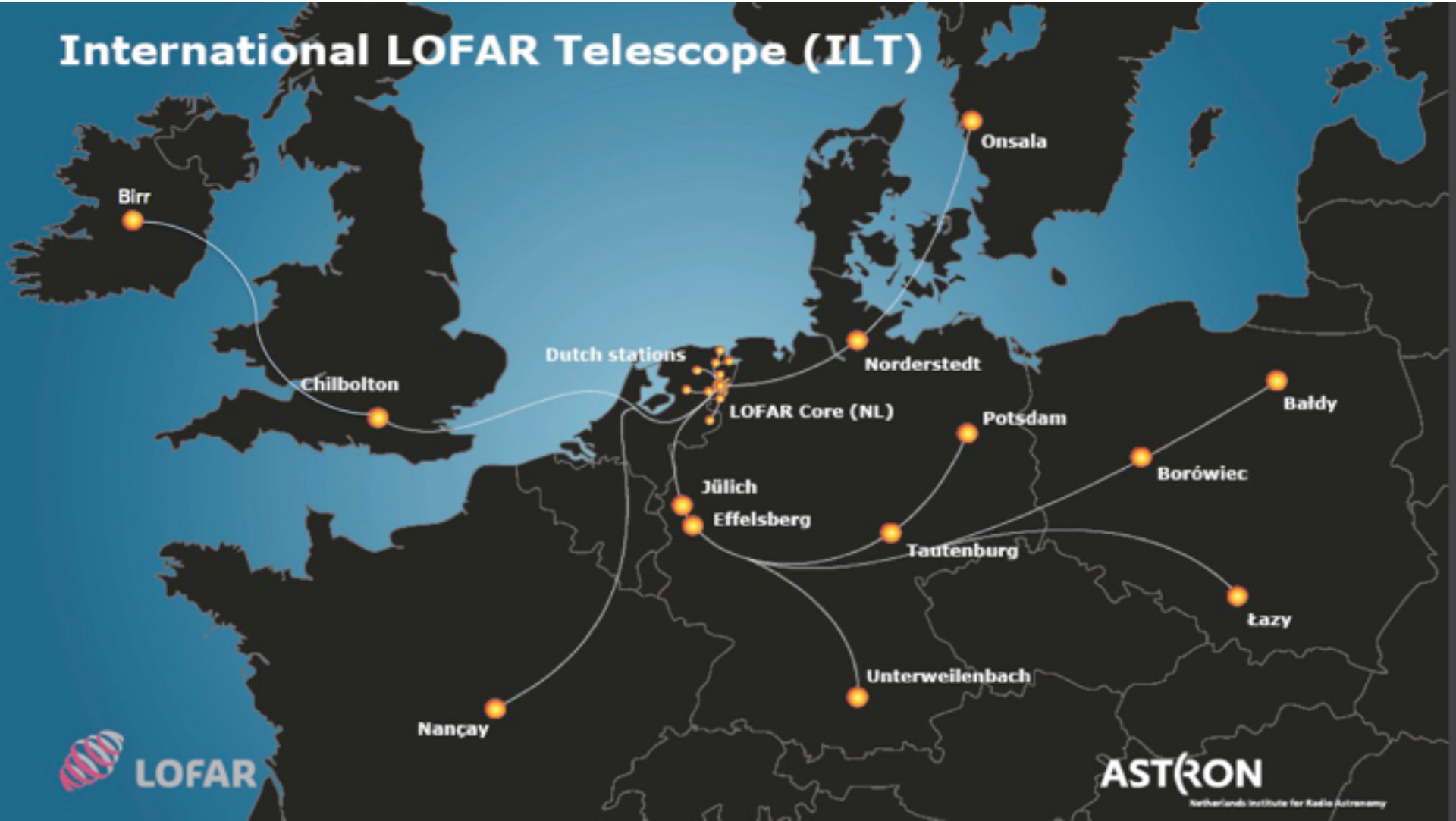
Power to the galaxies



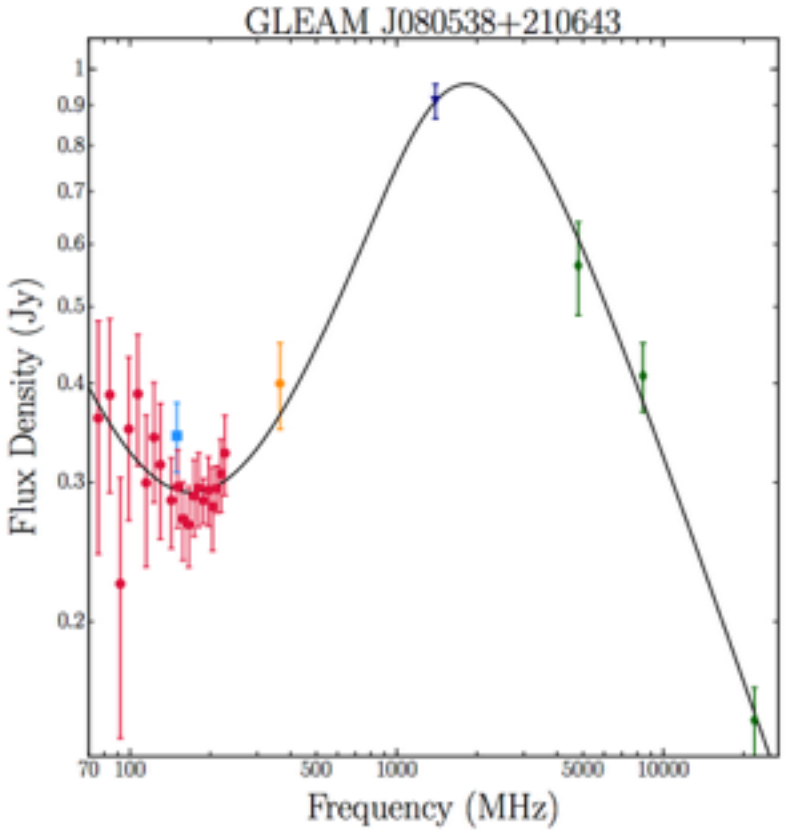
Callingham et al. (in prep.)

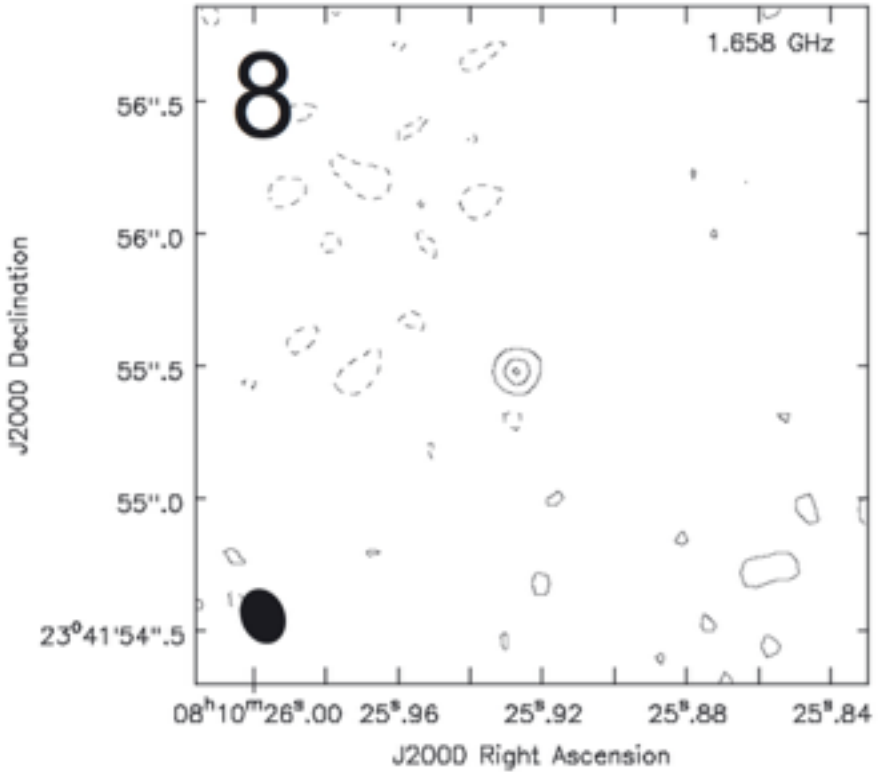
International baselines

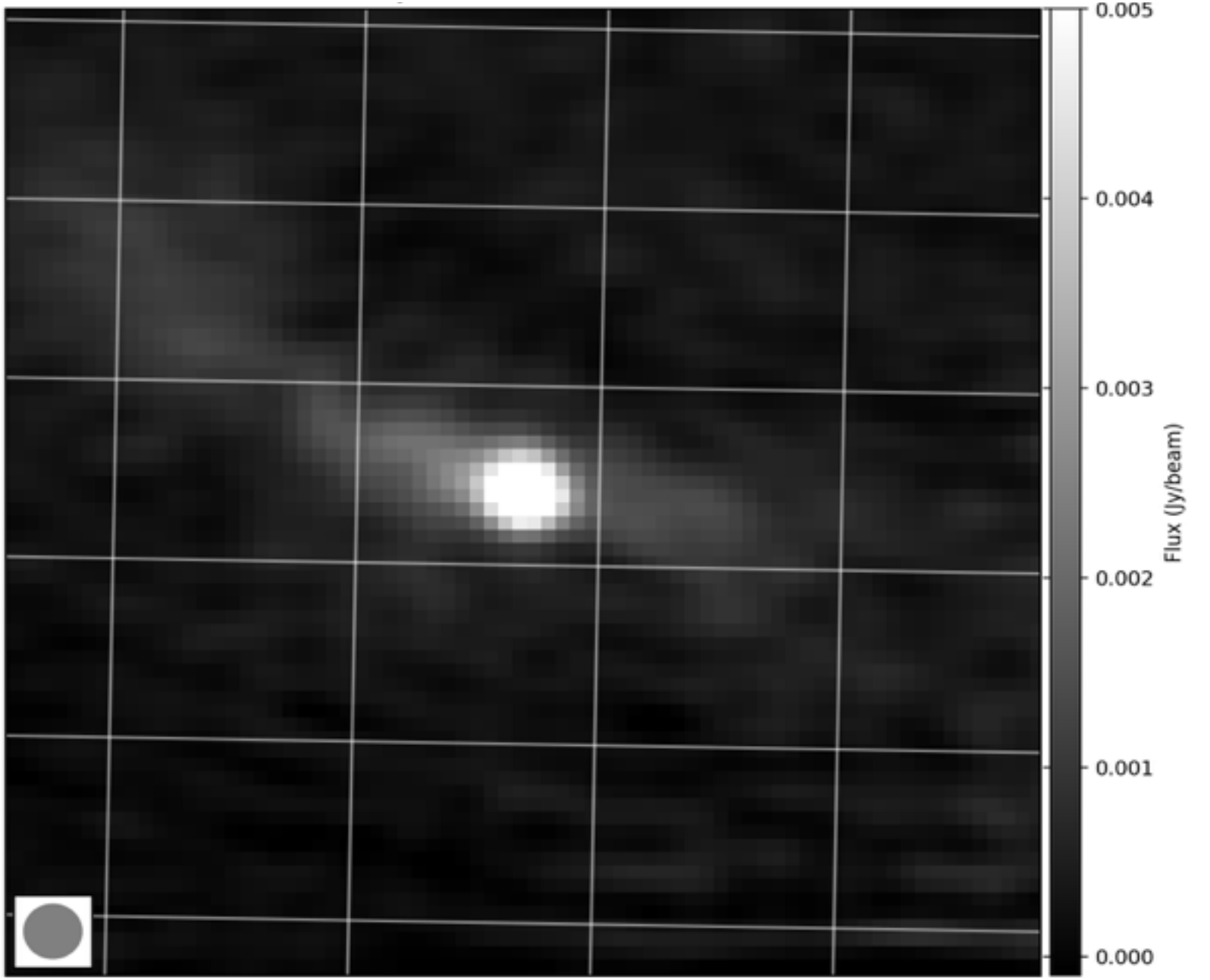
- LOFAR international baselines can achieve a resolution of ~0.8”.



Restarted?


GLEAM J080538+210643
Flux Density (Jy)
Frequency (MHz)
1
0.9
0.8
0.7
0.6
0.5
0.4
0.3
0.2
70
100
500
1000
5000
10000


8
1.658 GHz
56".5
56".0
55".5
55".0
23°41'54".5
J2000 Declination
08h10m26s.00
25s.96
25s.92
25s.88
25s.84
J2000 Right Ascension


0.005
0.004
0.003
0.002
0.001
0.000
Flux (Jy/beam)

Summary

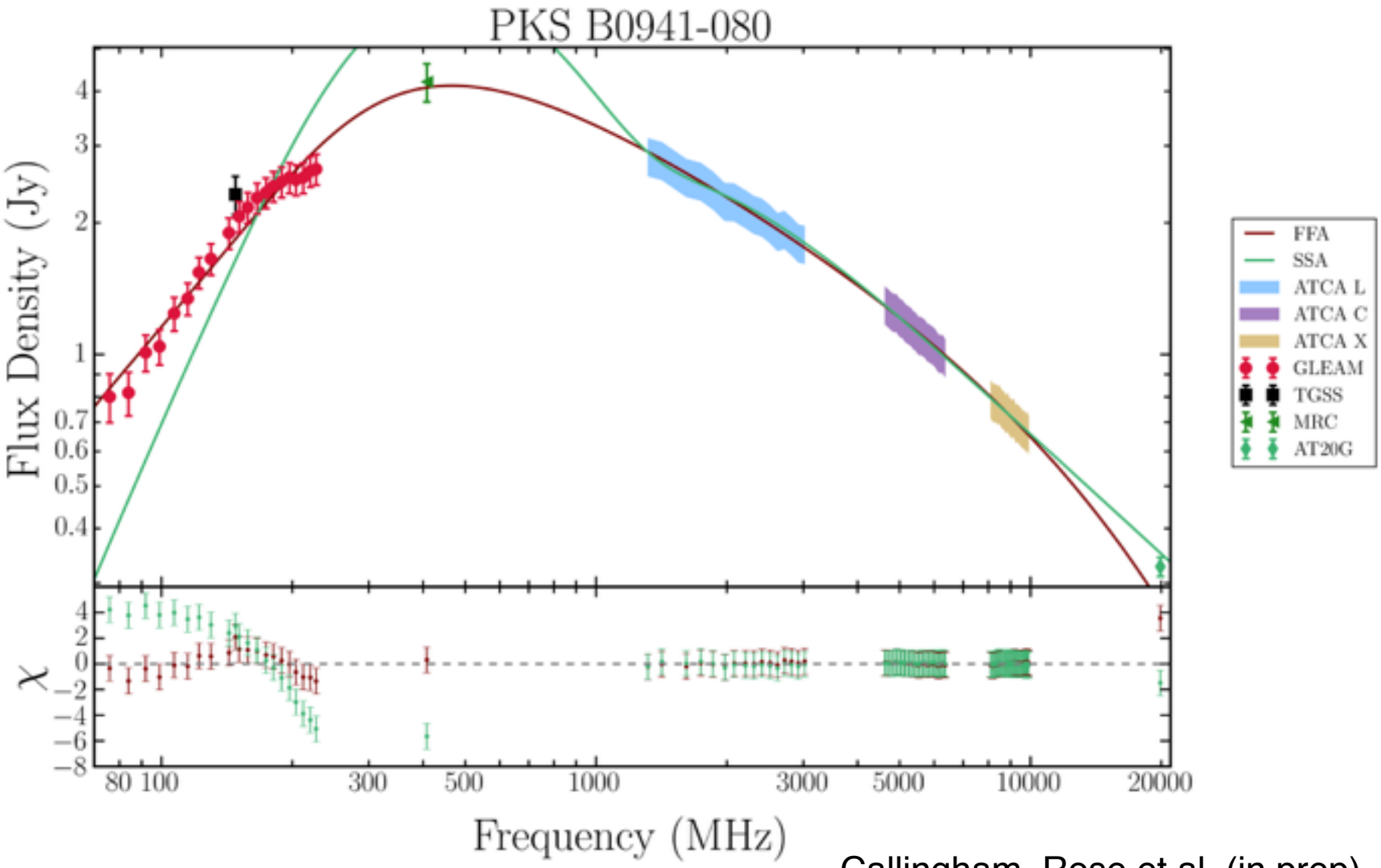
- Identified ~150 new peaked-spectrum sources in Hetdex field
- These sources are likely low luminosity counterparts to GPS sources identified at higher freq.
- Maybe dominant precursors to FR1 galaxies?
- Finding discrepancy in the number of GPS/CSS sources selected with LoTSS with those selected by GLEAM. Variability bias? Evolution?
- Using the spectra is a very useful way to find restarted GPS/CSS sources

MWA / Hurley-Walker

NASA, ESA, RIT, NRAO / AUI / NSF, Hubble Heritage

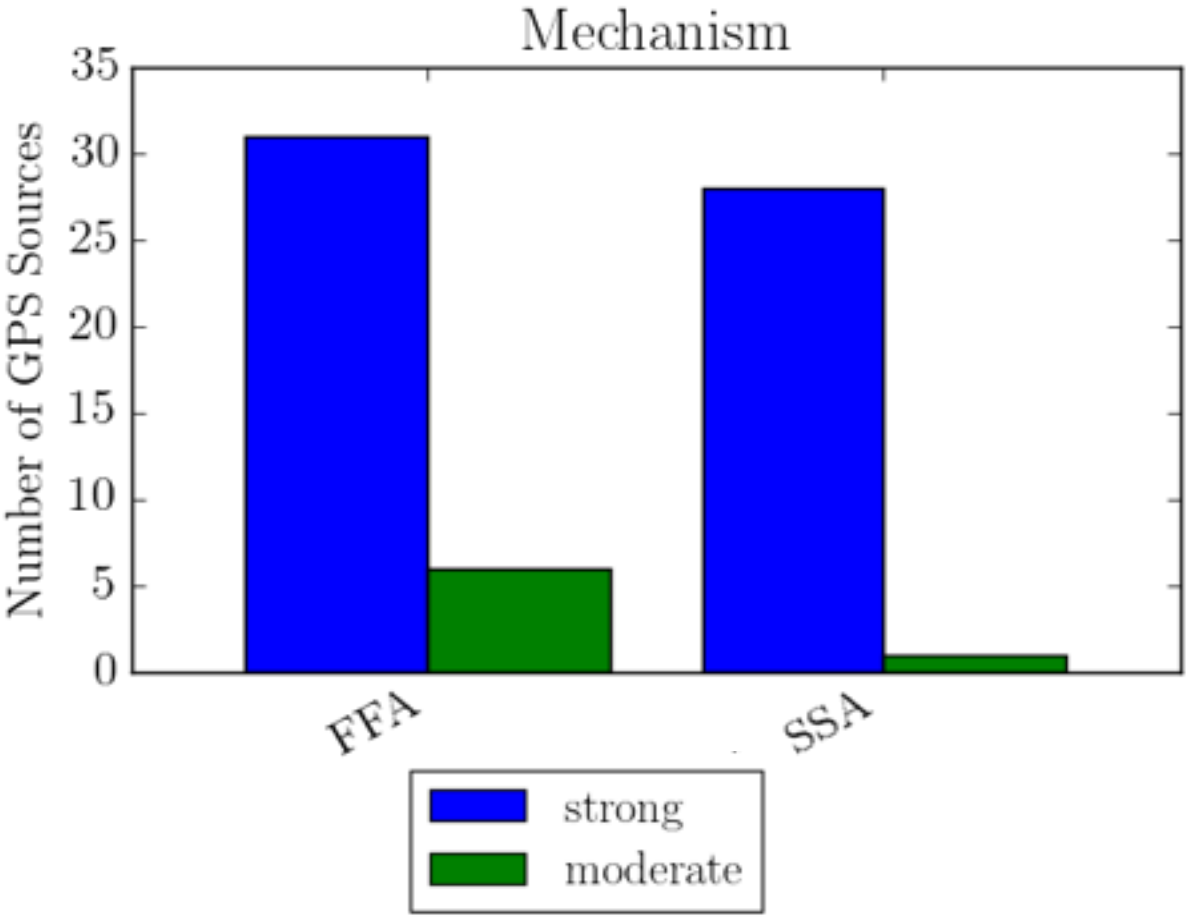
LOFAR / ASTRON

Bayesian Model Testing



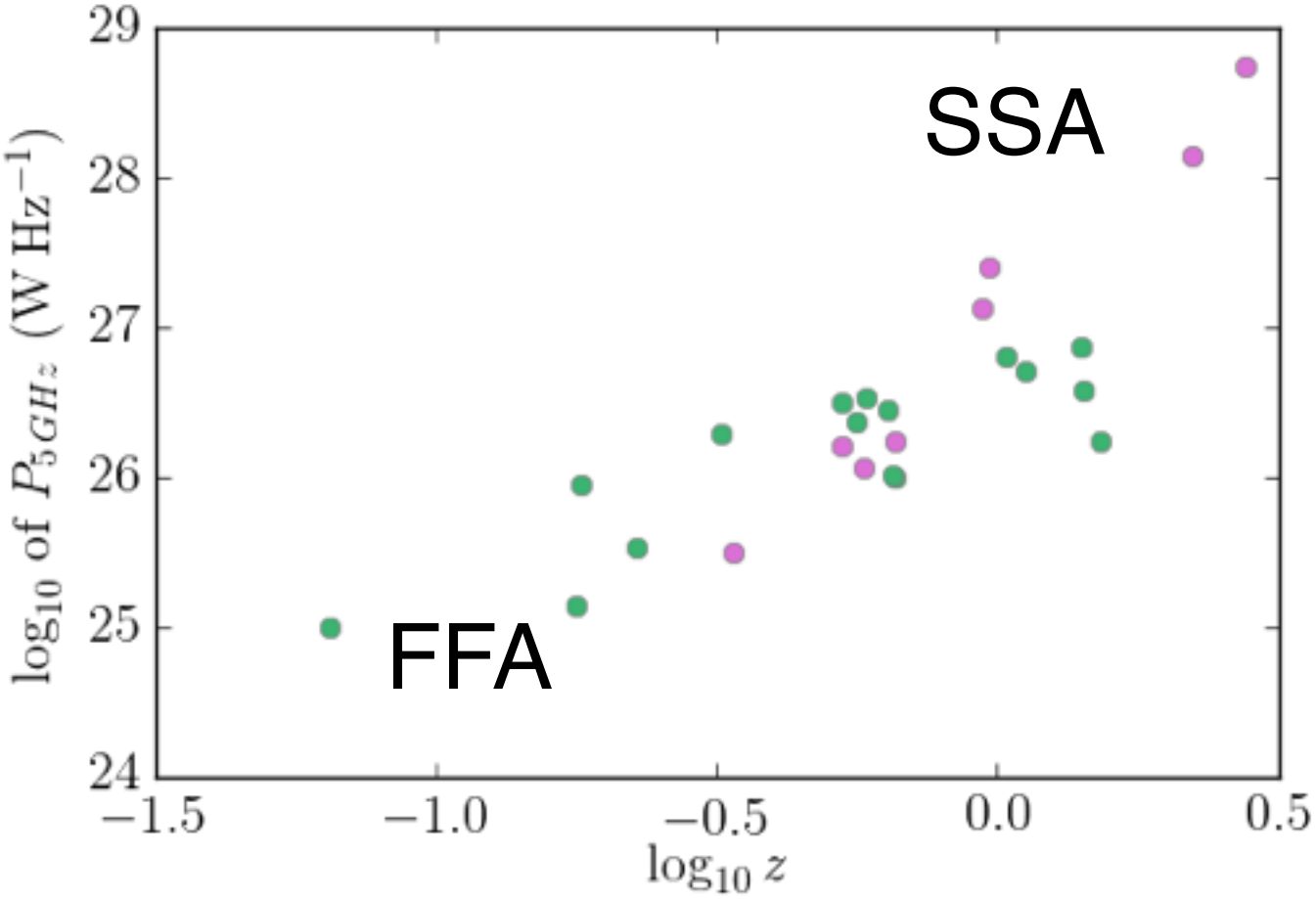
Callingham, Rose et al. (in prep).

Two populations!

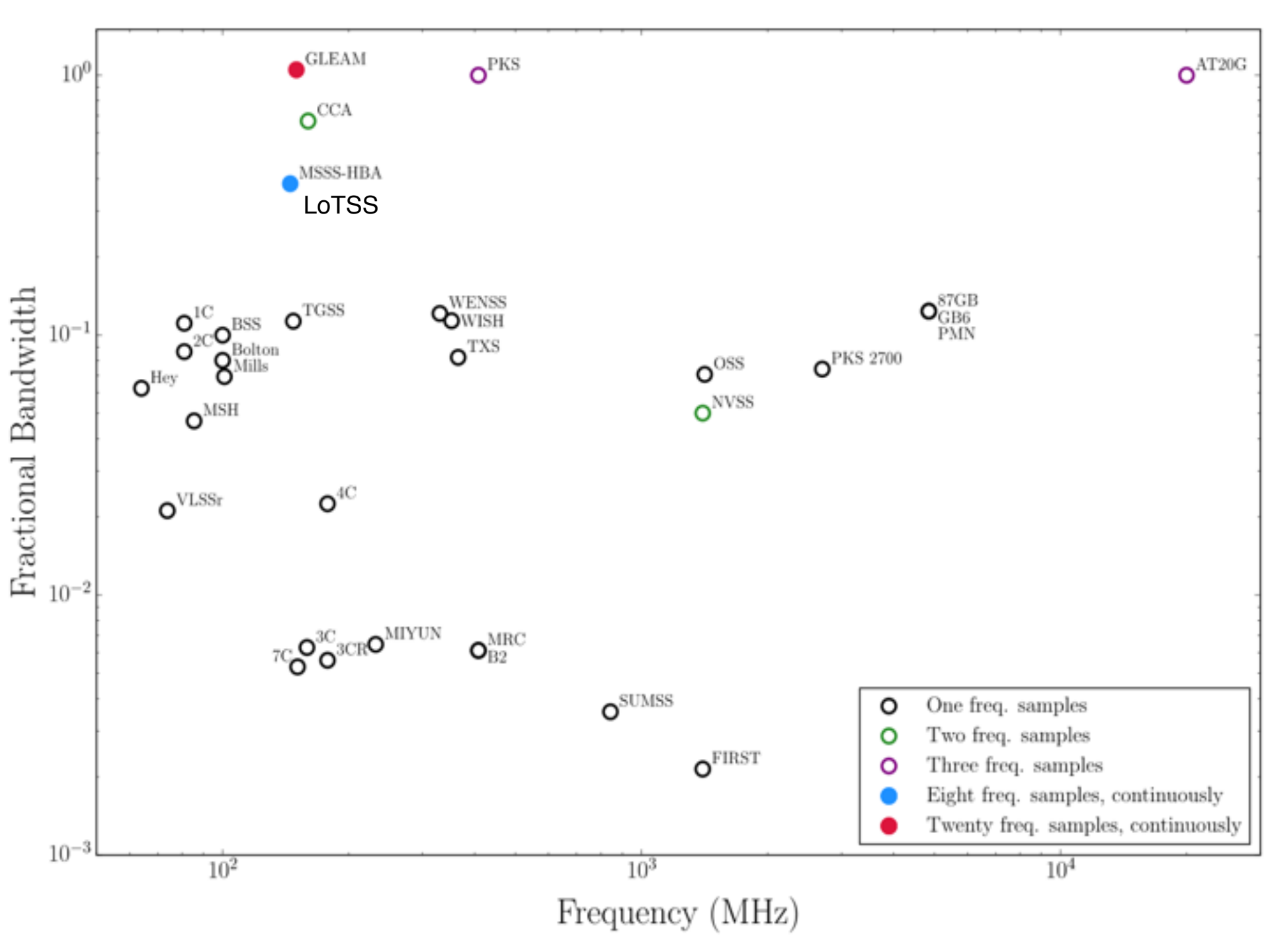


Callingham, Rose et al. (in prep).

Possible Power Dependence?

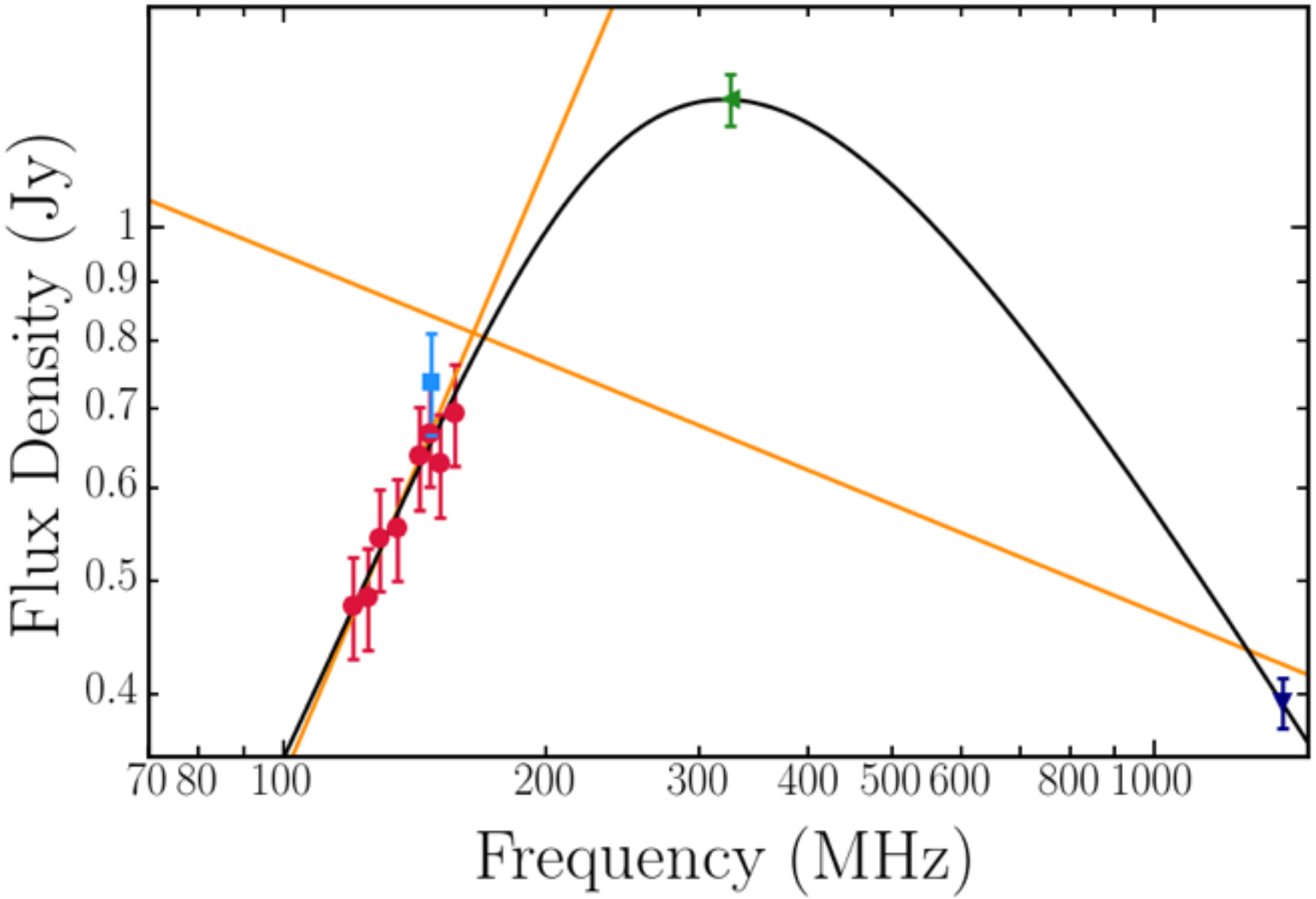


Callingham, Rose et al. (in prep).


GLEAM
PKS
AT20G
CCA
MSSS-HBA
LoTSS
1C
BSS
2C
Bolton
Mills
Hey
MSH
TGSS
WENSS
WISH
TXS
87GB
GB6
PMN
OSS
PKS 2700
NVSS
VLSSr
4C
7C
3C
3CR
MIYUN
MRC
B2
SUMSS
FIRST
One freq. samples
Two freq. samples
Three freq. samples
Eight freq. samples, continuously
Twenty freq. samples, continuously
Fractional Bandwidth
Frequency (MHz)
10^0
10^{-1}
10^{-2}
10^{-3}
10^2
10^3
10^4

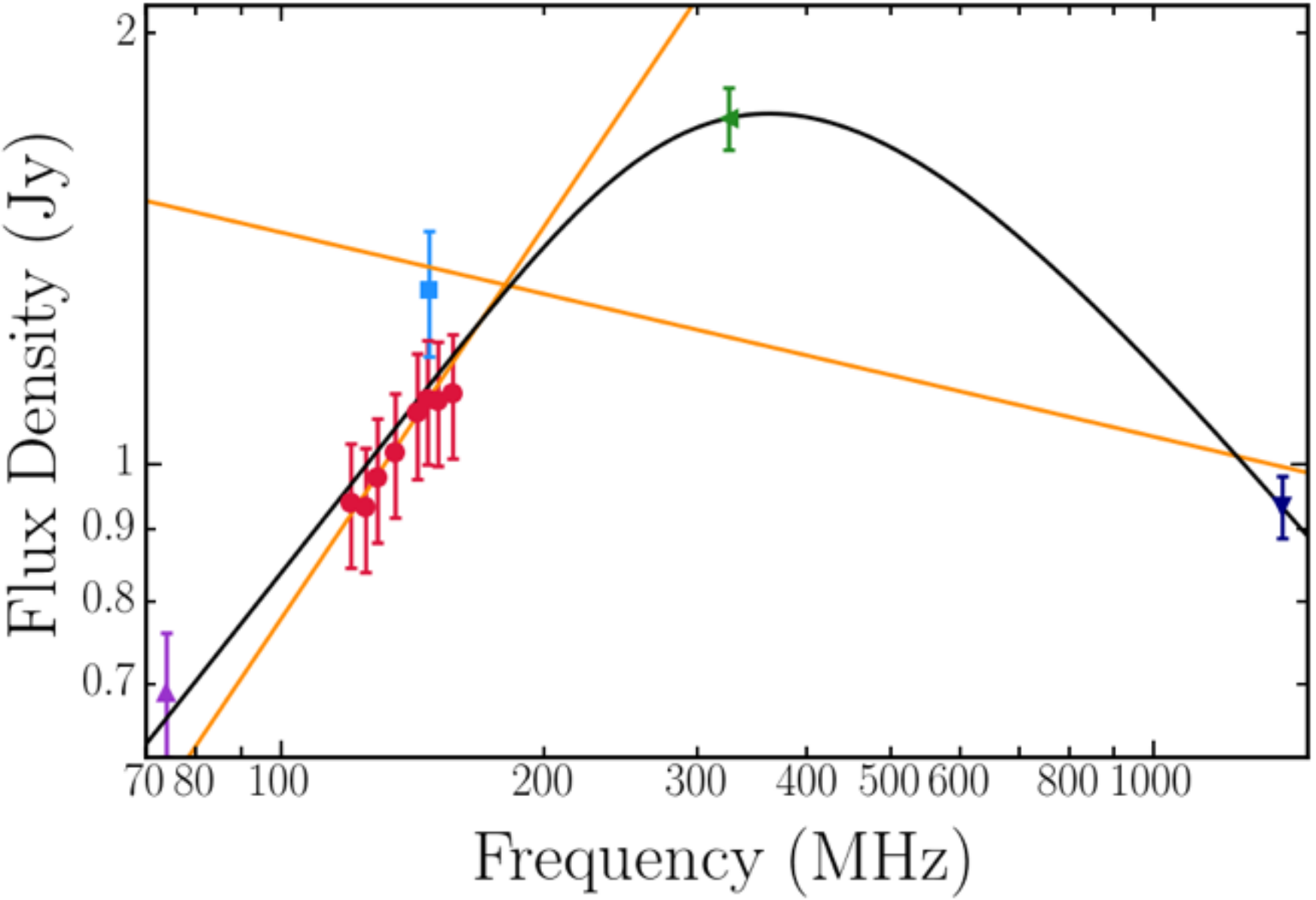
MSSS

ASTRON

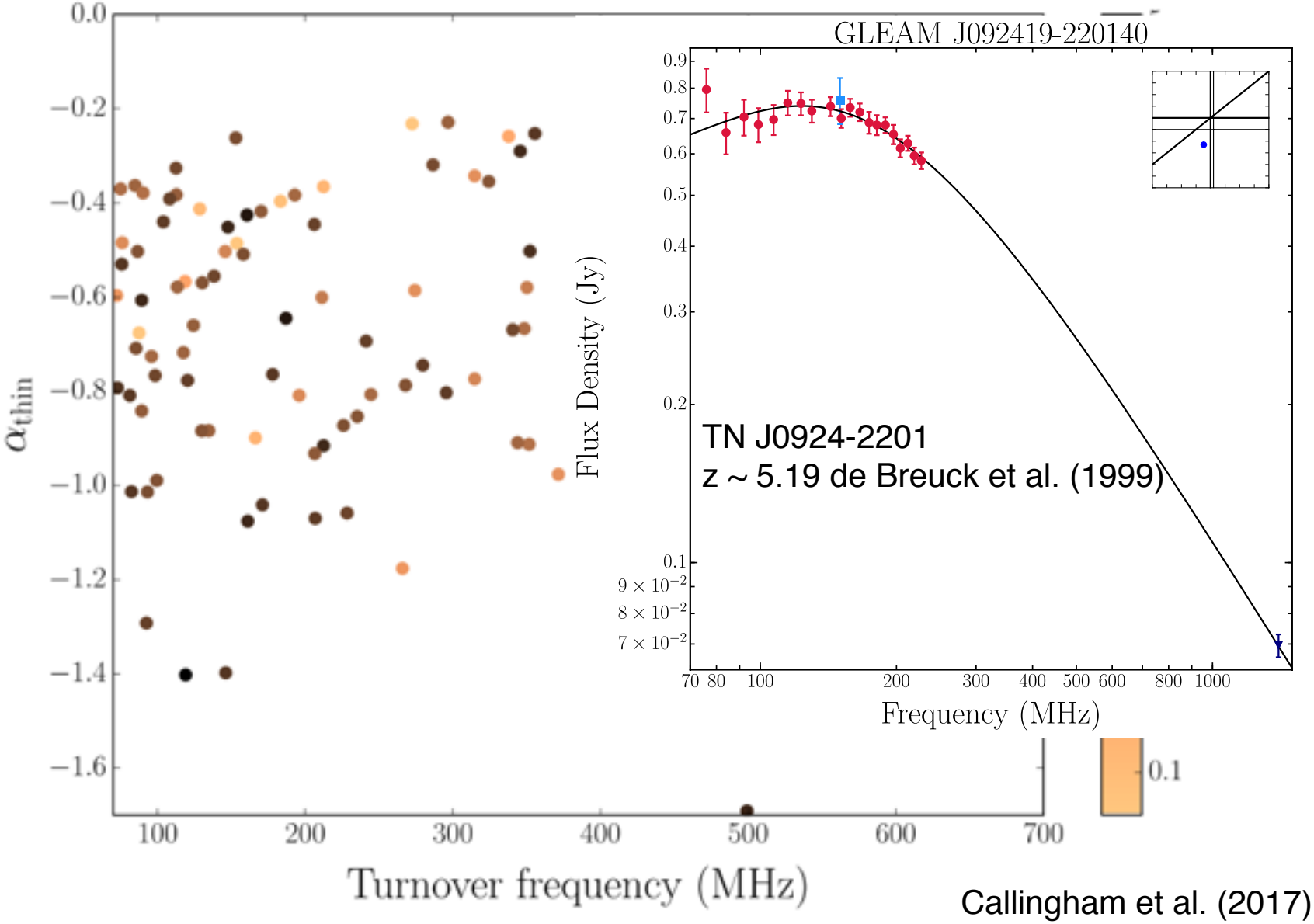

Flux Density (Jy)
1
0.9
0.8
0.7
0.6
0.5
0.4
70 80 100 200 300 400 500 600 800 1000
Frequency (MHz)

MSSS

ASTRON

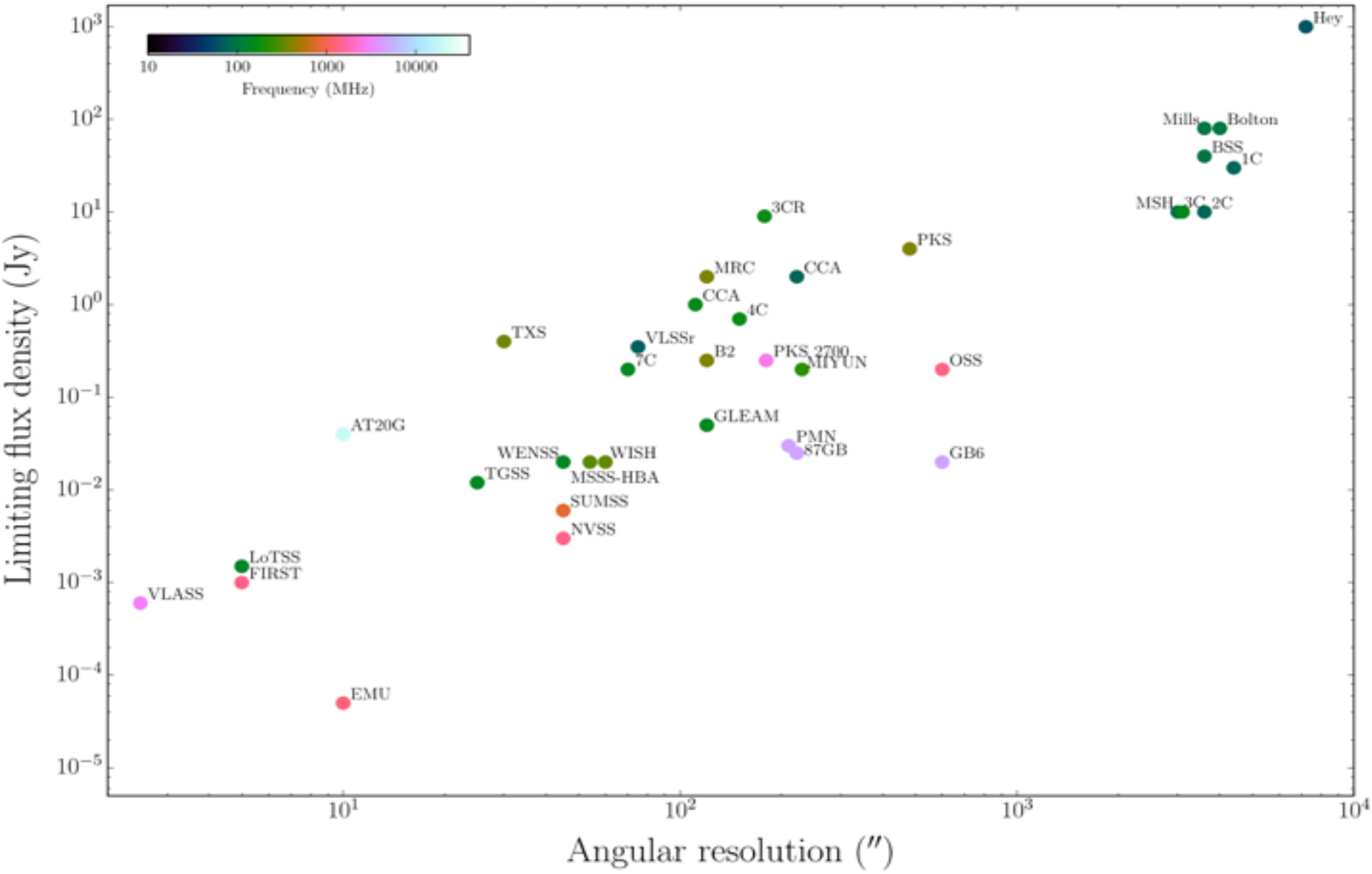

2
1
0.9
0.8
0.7
Flux Density (Jy)
70 80 100 200 300 400 500 600 800 1000
Frequency (MHz)

High Redshift Universe



Callingham et al. (2017)

What survey parameters make LoTSS look good?



Why Study GPS/CSS Sources?

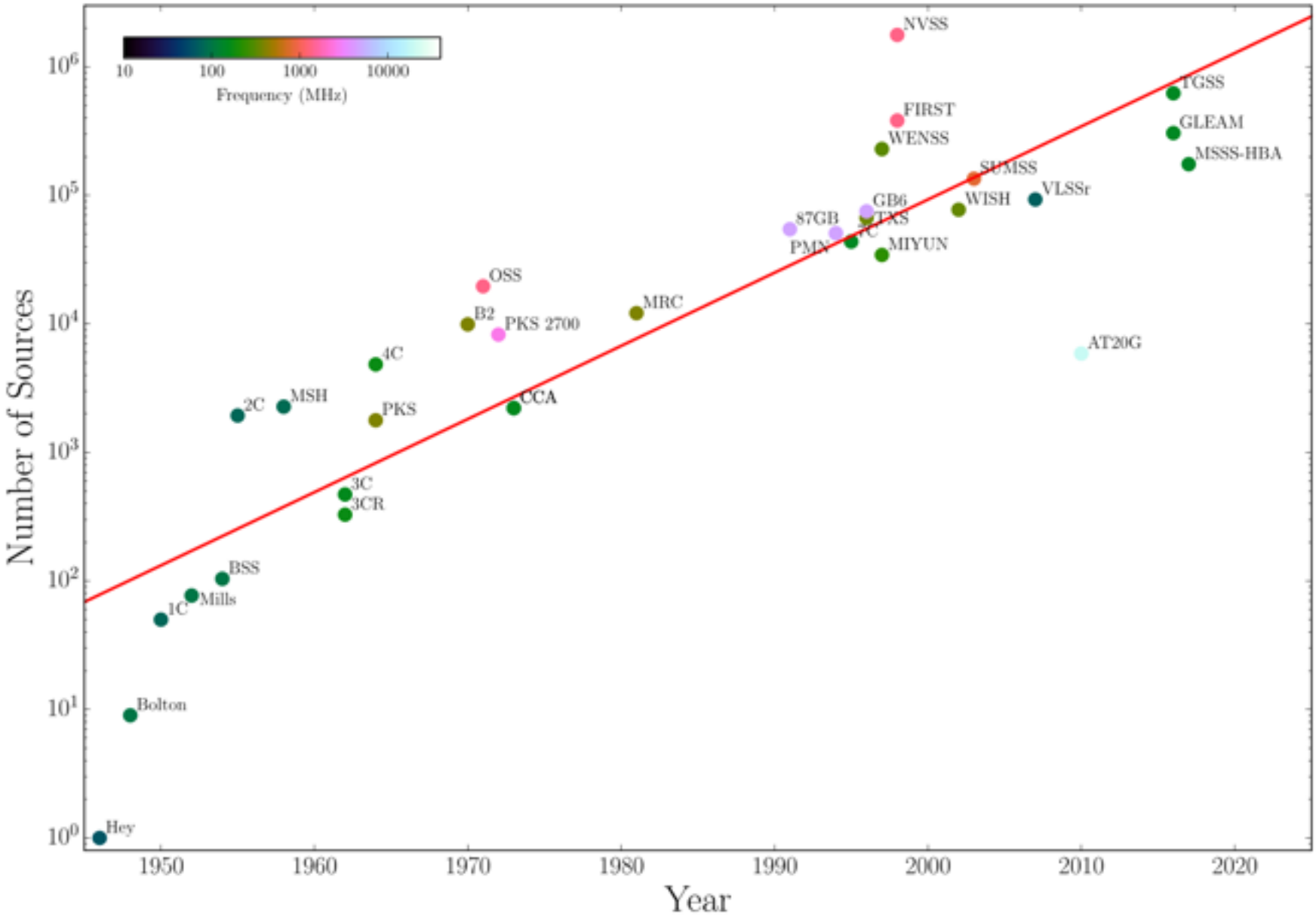
ASTRON

- Unique view of early AGN stages; probe of environment at scales of tens of pc
- Which radio galaxies evolve into "A team" sources (Cyg A, Her A, etc)?
- Are they confined to small spatial scales due to youth, frustration, or both?
- Cause of the turnover in spectrum? Free-free vs synchrotron self absorption

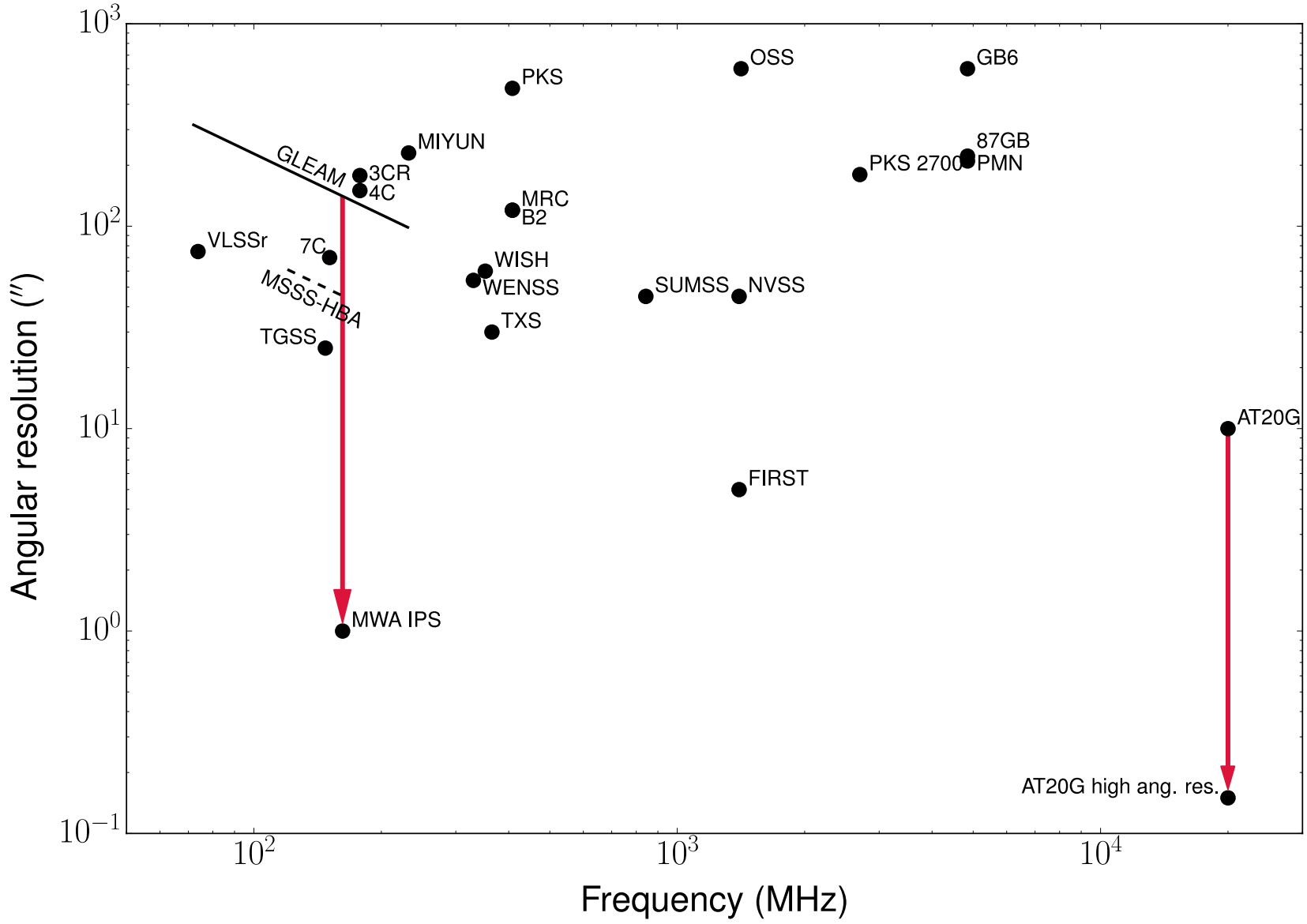
(see Peck et al. 1999; Kameno et al. 2000; Marr et al. 2001, 2014; Orienti & Dallacasa 2008; Tremblay et al. 2008; Tingay et al. 2015, Callingham et al. 2015)

NASA, ESA, S. Baum and C. O'Dea (RIT), R. Perley and W. Cotton (NRAO/AUI/NSF), and the Hubble Heritage Team

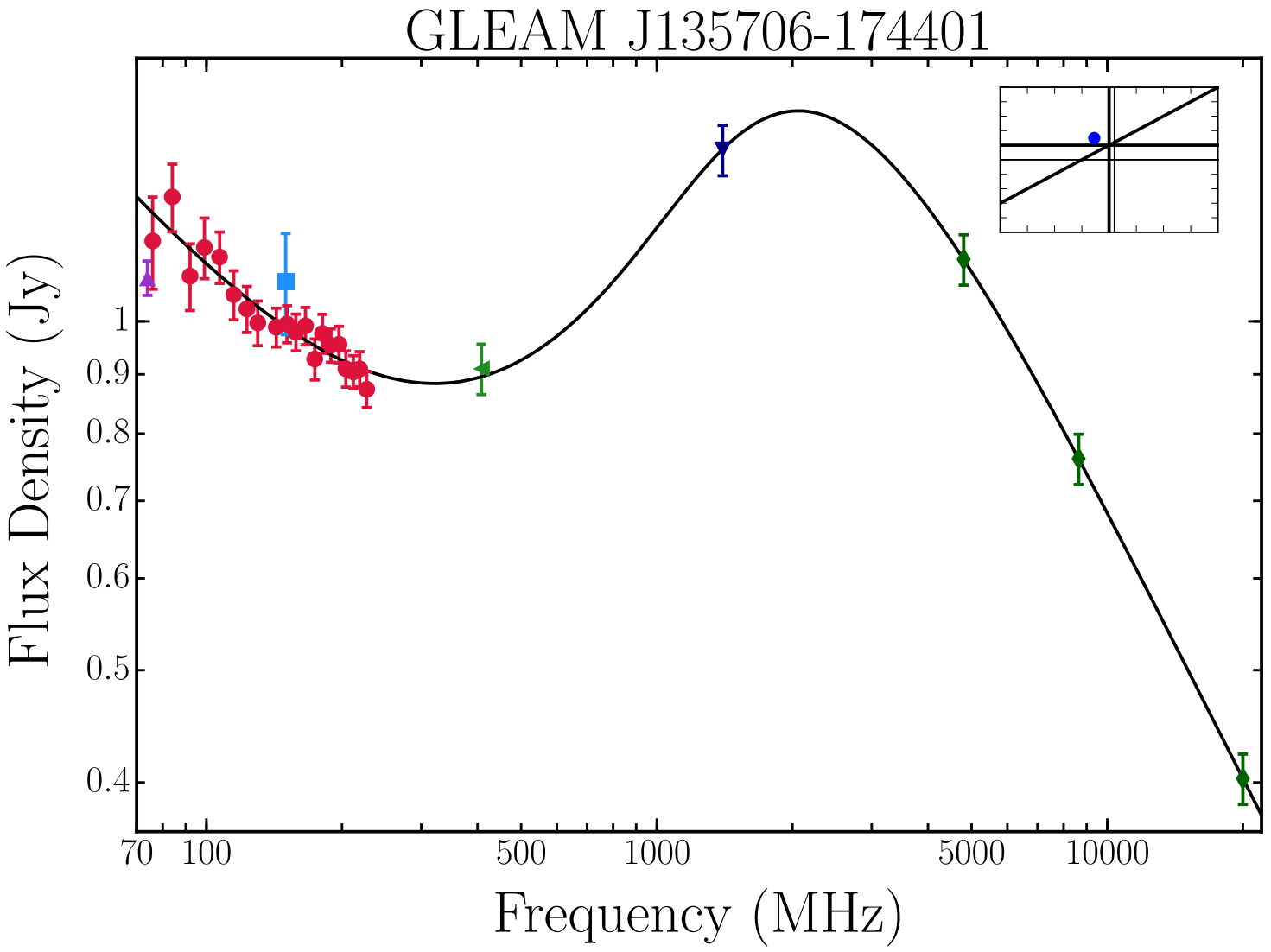
Widefield (continuum) survey evolution



IPS tricks


Angular resolution (″)
Frequency (MHz)
10^3
10^2
10^1
10^0
10^{-1}
10^2
10^3
10^4
PKS
OSS
GB6
MIYUN
GLEAM
3CR
4C
87GB
PKS 2700
PMN
MRC
B2
VLSSr
7C
WISH
WENSS
MSSS-HBA
SUMSS
NVSS
TXS
TGSS
AT20G
FIRST
MWA IPS
AT20G high ang. res.

Convex Source (II)



Callingham et al. (2017)